Long-term preservation of Hadean protocrust in Earth's mantle

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ABSTRACT

With plate tectonics operating on Earth, the preservation potential for mantle reservoirs from the Hadean Eon (>4.0 Ga) has been regarded as very small. However, many Archean rocks exhibit excesses of ^{182}W, the decay product of short-lived ^{182}Hf. The exact causes for these ^{182}W excesses, however, have remained ambiguous and it remains speculative, if the Archean ^{182}W anomalies and also ^{182}W deficits found in many young ocean island basalts (OIBs) mirror primordial Hadean mantle differentiation or just variable contributions from older meteorite building blocks delivered to the growing Earth. Here, we present a high precision ^{182}W isotope dataset for 3.22-3.55 Ga old rocks from the Kaapvaal Craton, Southern Africa. In expanding previous work, our study reveals widespread ^{182}W deficits in different rock units from the Kaapvaal Craton and also the very first discovery of a negative co-variation between short-lived ^{182}W and long-lived ^{176}Hf-^{143}Nd-^{138}Ce patterns, a trend of global significance. Amongst different models, these distinct patterns can be best explained by the presence of recycled mafic restites from Hadean protocrust in the ancient mantle beneath the Kaapvaal Craton. Further, the data provide unambiguous evidence for the operation of silicate differentiation processes on Earth during the lifetime of ^{182}Hf, i.e., the first 60 million years after solar system formation, thereby also providing lower bounds on the age of the Earth-Moon system. The striking isotopic similarity between recycled protocrust and the low ^{182}W endmember of modern OIBs might also be the missing link bridging ^{182}W isotope systematics in Archean and young mantle-derived rocks.

Due to plate tectonic processes, the accessible silicate reservoirs on Earth have lost most of their memory of the first ca. 500 Ma of Earth's history. Hence, our understanding of this time period comes from indirect evidence, e.g., from geochemical tracers such as short-lived and now extinct nuclide series that all were only active during the first ca. hundred million years after solar system formation(1–3). The detection of terrestrial variability in the relative abundances of short lived nuclide decay products such as ^{129}Xe, ^{142}Nd, and ^{182}W provided firm evidence that primordial reservoirs were not fully homogenized by mantle-dynamics, and played a significant role during the formation of the first continental crust(3–5). The recent discovery of ^{182}W, ^{142}Nd, and ^{129}Xe anomalies in modern mantle-derived rocks(2, 6, 7) demonstrates that ancient mantle reservoirs are still accessible. Whereas anomalous ^{129}Xe and ^{142}Nd isotope compositions in mantle-derived rocks can primarily be assigned to early planetary outgassing and early silicate differentiation, respectively, the presence of ^{182}W isotope anomalies can result from multiple processes. Negative ^{182}W anomalies in modern ocean island basalts (OIBs) were interpreted to result from core-mantle interaction(8, 9). In contrast, Archean rocks mainly exhibit elevated ^{182}W compositions. While some interpret the prevalence of positive ^{182}W anomalies in Archean rocks as a result of disproportional accretion(10), others have pointed out that this view may be a simplification as observations from other isotope systematics suggest other processes to be involved. Suggested alternative models invoke metal-silicate segregation or silicate differentiation in an early magma ocean, or during crust-mantle differentiation(11, 12) as explanations. Isotope anomalies of ^{142}Nd in Archean rocks clearly provide evidence for early silicate differentiation having operated during the Hadean, which potentially contributed to accompanying ^{182}W anomalies(12, 13). However, it has been demonstrated that pristine ^{142}Nd-^{182}W records are often obscured, either by multistage differentiation processes within the lifetime of ^{142}Nd, when ^{182}W was extant, or via fluid-controlled second stage metasomatic overprint of primordial ^{182}W patterns(14).

To further evaluate the processes that account for ^{182}W anomalies in Archean rocks, we investigated samples from the eastern Kaapvaal Craton, southern Africa. These lithologies are well suited to search for vestiges of early silicate differentiation, because they were shown to display both heterogeneous ^{142}Nd and ^{182}W compositions(11, 12, 15–17). We performed high-precision ^{182}W isotope analyses on a comprehensive suite of 17 samples that range from mantle-derived lithologies of mafic-ultramafic composition to different types of granitoids. These samples span an age range from ca. 3.55-3.22 Ga, represent the main lithological units of the Ancient Gneiss Complex (AGC) and also comprise the oldest rocks of the Barberton Granite-Greenstone terrain (BGGT). Moreover, most of these samples have previously been analyzed for their ^{143}Nd, ^{176}Hf, and ^{142}Nd compositions(15, 18, 19) and some samples were remeasured here as replicates. More information about the regional geology and samples is provided in SI Appendix. By combining ^{182}W isotope analysis with high-precision isotope dilution measurements for high field strength elements (HFSE), U, and Th, we assessed the sources of the W inventory in these samples. In order to place further constraints on Hadean processes, we performed ^{138}Ce isotope analyses to complement existing ^{143}Nd and ^{176}Hf data. Measurements of ^{182}W isotope compositions followed previously reported protocols(14, 20) that were slightly modified to yield sufficiently purified solutions for high-precision measurements using a Thermo

Fisher Neptune Plus MC-ICP-MS at Cologne. Uncertainties for averages of repeated analysis of sample solutions (95% confidence interval, n = 6-11) range between ±1.4 ppm and ±5.1 ppm (average ±2.7 ppm). Our intermediate precision is inferred from repeated analyses of in-house rock reference materials, also including a 3.27 Ga old komatiite from the Pilbara Craton (sample 160245, Ruth Well Formation), Western Australia, previously shown to display an excess of ^{182}W(20). All in-house rock reference materials were also passed through our separation protocol and measured in every session, yielding 2 SD between ± 1.5 ppm and ± 2.7 ppm (SI Appendix, Fig. S1). More information about the analytical protocol is provided in the method section.

Our results for ^{182}W isotope analysis are summarized in SI Appendix Table S1. Major and trace element compositions as well as ^{138}Ce-^{142,143}Nd-^{176}Hf isotope compositions are provided in SI Appendix, Table S2. Irrespective of petrology and provenance (AGC or BGGT), all rock types display ^{182}W isotope compositions that range from modern mantle values ($\mu^{182}W = 0$) to deficits as low as -9.2 ± 3.2 ppm. While most mantle-derived rocks from the BGGT display μ^{182}W values that overlap with the modern mantle value, most mantle-derived rocks from the AGC display resolvable μ^{182}W deficits. The distribution and the range of isotope compositions for ^{182}W in our rock samples from the Kaapvaal Craton is similar to that for ^{142}Nd, displaying both, negative and modern isotope composition(15). However, combined ^{182}W-^{142}Nd data for rocks from the eastern Kaapvaal Craton, also including literature data from the Schapenburg Greenstone Remnant (SGR) adjacent to the BGGT(12), only display a vague co-variation (SI Appendix, Fig. S2), even when only considering samples with pristine W concentrations (i.e., canonical W/Th ratios). Notably, our dataset reveals a negative co-variation of μ^{182}W with initial $\varepsilon^{143}Nd_{(t)}$ and $\varepsilon^{176}Hf_{(t)}$ for mantle-derived rocks (Fig. 1). To our knowledge, this is the first discovery of a co-variation between ^{182}W compositions and long-lived radiogenic nuclides. The observed co-variation for our samples is further strengthened by literature data for komatiites from the SGR adjacent to the BGGT(12) and the Komati Formation from the BGGT(11, 16).

A previous study(14) has shown that pristine ^{182}W isotope signatures can be modified during fluid-mediated second stage enrichment of W. One valuable tool to screen for disturbed elemental W budgets in mantle-derived rocks is the W/Th ratio, which displays a canonical range in pristine magmatic systems (0.09-0.24)(21). Although many mantle-derived rocks studied here reveal disturbed elemental W systematics (W/Th ≥ 0.24), they still display ^{182}W co-variations with initial $\varepsilon^{143}Nd_{(t)}$ and $\varepsilon^{176}Hf_{(t)}$ values. These observations indicate that the W redistribution was only of local character, in contrast to previous studies from Archean Cratons(14, 20). Moreover, co-variations with ^{182}W compositions are also observed for incompatible trace element ratios classically interpreted as immobile, in particular Hf/Sm and Zr/Sm (Figs. 2a+b). In addition, broadly coupled variations between the long-lived radiogenic isotopes and elemental co-variations with MgO (SI Appendix, Fig. S3) demonstrate that Hf and rare earth elements (incl. La, Ce, Sm, Nd, Lu) largely behaved immobile during metamorphism. Consequently, a metasomatic origin of the observed co-variations between ^{182}W and other radiogenic isotopes can be ruled out. As the elements involved display vastly different mobilities at metamorphic conditions, alteration would rather destroy the observed co-variations than forming them.

Most of the samples analyzed in this study reveal also strong correlations between their initial values of long-lived radiogenic isotopes like $\varepsilon^{143}Nd_{(t)}$, $\varepsilon^{176}Hf_{(t)}$ and $\varepsilon^{138}Ce_{(t)}$ (Figs. 3a+b). Only two outliers (AGC 38 and ZA-38) display disturbed initial $\varepsilon^{138}Ce_{(t)}$ values but still preserve pristine $\varepsilon^{143}Nd_{(t)}$ and $\varepsilon^{176}Hf_{(t)}$ systematics. Therefore, initial $\varepsilon^{138}Ce_{(t)}$ values for AGC 38 and ZA-38 are excluded from further interpretations and are not shown in Fig. 3b. In this regard, combined ^{143}Nd-^{176}Hf-^{138}Ce systematics serve as a valuable tool to clarify why two mantle-derived rocks (AGC 350 and ZA-31a, pale red symbols, Fig. 1) slightly deviate from the $\mu^{182}W$ vs. $\varepsilon^{143}Nd_{(t)}$ and $\varepsilon^{176}Hf_{(t)}$ trends, most likely indicating slight secondary enrichment by somewhat more negative $\mu^{182}W$. As the observed co-variations of $\mu^{182}W$ with $\varepsilon^{143}Nd_{(t)}$ and $\varepsilon^{176}Hf_{(t)}$ are defined by mafic-ultramafic volcanic rocks, it is obvious that the observed trend reflects mixing between different mantle-source reservoirs. One mantle endmember exhibits no resolvable ^{182}W isotope anomalies at near chondritic initial $\varepsilon^{143}Nd_{(t)}$ and $\varepsilon^{176}Hf_{(t)}$ values, most likely representing near primitive mantle. The other endmember is best characterized by komatiites from the SGR that exhibit the largest ^{182}W isotope deficits extending to -11.4 ppm and strongly elevated initial $\varepsilon^{143}Nd_{(t)}$ and $\varepsilon^{176}Hf_{(t)}$ values of up to +2.6 and +6.2, respectively(12). It is surprising that felsic samples from the Kaapvaal Craton plot on the same trend as mafic samples, suggesting short residence times between emplacement of the mafic protolith and formation of felsic orthogneisses (open symbols SI Appendix, Fig. S4).

In the following discussion, we will largely focus on the origin of the low ^{182}W endmember. As we will show, the low ^{182}W endmember may provide novel insights into the secular evolution of the ^{182}W isotope composition of Earth´s mantle. In particular, we evaluate, if present-day mantle plumes with their characteristic ^{182}W deficit may be modern analogues of the low ^{182}W endmember from the Kaapvaal Craton. The presence of ^{182}W deficits could be the consequence of several processes. These include (1) equilibration of the mantle source with anomalously large amounts of late accretionary material (late accretion hypothesis), (2) early fractionation of Hf from W by silicate crystal-liquid fractionation, e.g., in an early magma ocean, or (3) core-mantle interaction.

The late accretion hypothesis has been postulated to explain the relative and absolute abundances of highly siderophile elements (HSE) in the bulk silicate Earth (BSE) by the addition of about 0.5% of chondritic material after core formation(22, 23). Late accretion would not only have affected the HSE budget of the BSE but also its ^{182}W isotope composition(3). Accordingly, some portions of the Archean mantle could have remained in disequilibrium(24), and mantle domains that did not fully equilibrate with late accretionary components, would be characterized by positive ^{182}W isotope anomalies and HSE abundances that are lower than the modern BSE. Consequently, negative ^{182}W isotope anomalies would imply excesses of late accreted components that would also be reflected in unusually high HSE contents. The absolute HSE abundances in the mantle source of the SGR-like endmember with its large ^{182}W deficit, however, were estimated to be only ca. 30% of those in the present-day BSE(12, 25). Accordingly, a late accretion model can therefore be ruled out. Direct addition of core material to the source of the SGR komatiites also seems an unlikely explanation, as this would likely result in more elevated PGE concentrations(26),

again, in contrast with the low absolute HSE abundances in the source of the SGR-like endmember(25).

An alternative explanation for negative ^{182}W isotope anomalies in Archean rocks like those from the Kaapvaal Craton may be offered by insights from recent studies on OIBs. It has been proposed that prevalent negative ^{182}W isotope anomalies in modern, plume-derived OIBs result from chemical and isotopic equilibration between their mantle sources and the outer core(8, 27). Alternatively, noble gas work on modern mantle-derived rocks suggested that the source reservoirs must have had differentiated from the convecting mantle very early (pre 4.45 Ga)(5, 28, 29). The concurrent ^{182}W isotope anomalies in the modern mantle may therefore also have been produced early by in-situ decay of ^{182}Hf (i.e., during the first ca. 60 Ma after solar system formation). The presence of such ancient mantle reservoirs and the role of mantle plumes in the past, in particular their contribution to the secular evolution of the ^{182}W isotope composition in the BSE has so far only poorly been constrained. Notably, it has been argued that Archean mafic-ultramafic sequences like those in the BGGT also originate from mantle plumes(30). In this regard, mantle derived rocks from the Kaapvaal Craton may have preserved vestiges of ancient mantle heterogeneities, similar to young OIBs.

It has been postulated that recycling of crustal material is responsible for the geochemical and isotopic variability of modern mantle plumes(31). In the case of the Kaapvaal Craton, direct recycling of ancient protocrust formed during the first ca. 60 Ma appears unlikely, because in this case the negative ^{182}W and ^{142}Nd anomalies should be coupled with unradiogenic ^{143}Nd and ^{176}Hf compositions. In this regard, the coupled depletions of ^{182}W and ^{142}Nd and Hf-Nd isotope patterns led previous studies(12, 17) to conclude that the komatiites from the SGR derived from a mantle domain that was enriched very early (ca. 30 Ma after solar system formation) in highly incompatible elements as a result of fractionating a Mg- and Ca-perovskite mineral assemblage in an early magma ocean. When originally proposed(12) this conclusion was mainly based on apparently decoupled initial $\varepsilon^{143}Nd_{(t)}$ and $\varepsilon^{176}Hf_{(t)}$ compositions of BGGT rocks(16, 32) (see grey symbols in SI Appendix Fig. S5). However, more recent work re-investigated mafic-ultramafic samples from the BGGT (Komati, Sandspruit, Theespruit Formations)(18) and AGC (Dwalile Greenstone Remnant)(19) by employing more sophisticated sample dissolution protocols, yielding considerably less scatter (see red symbols in SI Appendix Fig. S5). In fact, together with komatiites from the SGR(12), the new data fall on a trend closely resembling the modern mantle array(33), in line with the consideration that the terrestrial Hf-Nd mantle array has already been established on the early Earth(34) (Fig. 3a). On the basis of more rigorous modelling, employing updated sets of partition coefficients(35, 36) and by adopting a previous model for SGR komatiites(12), we re-evaluated the control of perovskite segregation and subsequent mantle depletion on the ^{143}Nd and ^{176}Hf isotope inventory (blue symbols, Fig. 3a). The modelled ^{143}Nd and ^{176}Hf isotope compositions for perovskite segregation are not in accord with observations from mafic ultra-mafic rocks of the Kaapvaal Craton, and also not with those of the SGR komatiite suite.

In line with our results for ^{143}Nd-^{176}Hf, initial $\varepsilon^{138}Ce$ and $\varepsilon^{143}Nd$ systematics of our samples also closely fall on the modern terrestrial array(37) and do not match the

modeled compositions of magma ocean relics (blue symbols, Fig. 3b). A recent study on Archean rocks from the Pilbara Craton has introduced the $^{138}La^{138}Ce$ system as a novel tool for constraining Archean mantle evolution and in particular the role of perovskite fractionation in an early magma ocean(38). Assuming that perovskite fractionation affected the parental sources of mantle-derived rocks from the Kaapvaal Craton it would be expected that initial $\varepsilon^{138}Ce$ compositions would be much more unradiogenic (Fig. 3b). On the basis of available literature data, we cannot rule out that other mafic units in the BGGT (e.g. Weltevreden, see blue field in SI Appendix Fig. S5) do indeed preserve an extreme decoupling in their $\varepsilon^{143}Nd$ and $\varepsilon^{176}Hf$ systematics and anomalous $\varepsilon^{138}Ce$. Notably, however, a recent study on the 3.33 Ga Commondale komatiites from the Kaapvaal Craton demonstrates that decoupled ^{143}Nd-^{176}Hf patterns are not unique to magma ocean relics but can also be generated via melting events in the stability field of garnet combined with the involvement of residual garnet-bearing pyroxenites(39). In summary, our combined ^{143}Nd-^{176}Hf-^{138}Ce data can now show that there is no clear evidence for reservoirs that underwent perovskite fractionation in an early magma ocean. A more detailed assessment of early magma ocean processes involving perovskite segregation is given in the method section.

Based on the considerations above, a two-stage process is clearly required where the negative ^{182}W and ^{142}Nd anomalies formed early and the radiogenic ^{143}Nd and ^{176}Hf compositions were established after the short-lived systems went extinct. Our preferred geodynamic model is illustrated in Fig. 4 and a detailed description is given in the method section. Our model is inspired by previous studies on the formation mechanisms of early continental crust(40, 41). Accordingly, after formation of mafic protocrust (Fig. 4a) intra-crustal fractionation lead to the formation of a felsic, TTG-like crust and mafic lower crustal restites that are recycled into the upper mantle (Fig. 4b). This process can account for hybrid mantle reservoirs, where ambient upper mantle mixes with mafic restites from the lower crust (Fig. 4c). Melting of such hybrid reservoirs might have been triggered by ascending plume-like material from the lower mantle that may also account for the compositional trend observed for mafic rocks from the Kaapvaal Craton. In this scenario, the lower mantle endmember is characterized by the Barberton komatiites and the hybrid upper mantle endmember is characterized by the SGR komatiites. As demonstrated below, such mixing relationships provide a plausible explanation for the negative co-variation between short- and long-lived radiogenic systems. Moreover, our model can also explain the incompatible trace element systematics in our samples and the SGR komatiites (SI Appendix, Figs. S3 and S6).

Following constraints from phase equilibrium and trace element modeling, melting of Archean TTG suites from mafic protocrust leaves behind residual assemblages of amphibolitic, garnet-amphibolitic or garnet-pyroxenitic composition(42). Figure 4a illustrates that during a first stage mafic protocrust that formed 50 Ma after solar system formation developed strongly unradiogenic isotope compositions, in particular for the short-lived decay products ^{182}W and ^{142}Nd. Subsequent TTG melting (stage 2 in Fig 4a) leaves behind garnet-rich restites(43), and depending on the timing of this second event, the residual restites will develop towards markedly different ^{142}Nd isotope compositions with time. In contrast, the ^{182}W isotope composition will be insensitive to the timing of TTG extraction, because ^{182}Hf went extinct shortly after formation of the

protocrust. Evidence for the presence of such ancient TTG precursors in the Kaapvaal Craton comes from Hf-in-zircon isotope data(44–46) and from rare Hadean detrital zircons from the ca. 3.3 Ga Fig Tree Formation(47) that suggest formation of a felsic protocrust already by the Eoarchean or late Hadean. Due to the longer half-lifes of their parent nuclides, ^{143}Nd and ^{176}Hf isotope compositions in the restites integrate a larger time span and develop much less heterogeneous with time than ^{142}Nd, which can only be formed over a smaller time interval until ^{146}Sm becomes extinct. These considerations explain, why ^{142}Nd signatures became quite variable, depending on the time of TTG extraction, unlike long-lived Hf-Nd compositions that persistently developed towards slightly radiogenic values over time.

It becomes apparent from Fig. 5a that lower crustal restites from ancient protocrust can explain why the ^{143}Nd and ^{176}Hf isotope compositions are so tightly correlated with ^{182}W but not with ^{142}Nd. Recycling of such restites into the upper mantle formed a hybrid source that is best approximated by compositions of Schapenburg komatiites, which formed through high-degree melting. We found that 10-20% of restites admixed to an ambient depleted upper mantle already reproduce the radiogenic isotope compositions found in the SGR endmember (Fig. 5b). Once this hybrid upper mantle source is mixed with primitive material supplied by ascending mantle plumes, it can account for the systematic coupling between initial $\varepsilon^{143}Nd_{(t)}$- $\varepsilon^{176}Hf_{(t)}$ and $\varepsilon^{143}Nd_{(t)}$- $\varepsilon^{138}Ce_{(t)}$ (Fig. 3) and also for the opposing variations of ^{182}W (Fig. 1). Exact modeling of ^{138}La-^{138}Ce systematics is hampered by their poorly constrained behavior during mantle melting, where La-Ce behave highly incompatible and modelled La/Ce is extremely dependent on melt porosity. Notably, our proposed model can well reproduce the incompatible trace element compositions and reconcile distinct trace element features that are diagnostic for the SGR komatiites. As shown in SI Appendix Fig. S6 our modeling results are in good agreement with the SGR komatiites originating from 20-30% batch melting of a hybrid source that consists of ambient depleted mantle and 10-20% lower crustal restites. Moreover, our model can reproduce distinctive Hf/Sm and Zr/Sm ratios prominent within SGR komatiites and their co-variation with ^{182}W isotope compositions (Figs. 2a+b). A recent study on mantle derived rocks from the Kaapvaal Craton(17) reported a similar correlation of Hf/Sm with ^{142}Nd compositions arguing that this feature is unique to deep magma ocean crystallization processes that happened soon after Earth accretion. Our lower crustal restite model can now offer an alternative explanation. The hybrid source model can also explain the Re-Os isotope inventory of the SGR komatiites(25). As shown in the method section the positive initial γ^{187}Os of the SGR komatiites (γ^{187}Os = +3.7 ± 0.3)(25) is in accord with the addition of 10-20 % restites from Hadean protocrust to a depleted mantle source that was previously exhausted in sulfides.

The lower crustal restite model for the Kaapvaal Craton presented in this study also provides an intriguing explanation of ^{182}W isotope variations in modern OIBs. A recent study(48) has proposed that the global ^{182}W dataset for OIBs can be explained by the admixture of the classical mantle endmember components DMM (depleted MORB mantle), EM1 (enriched mantle I), EM2 (enriched mantle II), and HIMU (high “μ” or ^{238}U/^{204}Pb) to a primordial reservoir that is characterized by negative ^{182}W anomalies and depleted ^{143}Nd/^{144}Nd composition. Remarkably, lower crustal restites, formed between 4.35 and 4.25 Ga from ancient mafic protocrust constitute a viable

endmember for the global OIB array in $\mu^{182}W$-$^{143}Nd/^{144}Nd$ space (references(8, 48) and refs. therein), once calculated to present day $^{143}Nd/^{144}Nd$ (Fig. 6). We therefore speculate, that lower crustal restites from Hadean protocrust were delaminated and ultimately recycled into the lower mantle where they might have become part of large low shear-wave velocity provinces (LLSVPs) in the present day mantle that are interpreted to contribute to rising mantle plumes(49). Indeed, it has been shown that the modeled restite has the potential to delaminate into the mantle due to its density contrast compared to ambient mantle(43). Accordingly, geophysical studies demonstrated that LLSVPs may represent mixtures of recycled dense material that accumulated at the core mantle boundary(50, 51). Taking into consideration that the strongly depleted restites would exhibit low He abundances, recycling of this component into the lower mantle would not significantly affect the $^{3}He/^{4}He$ ratios of a primordial undegassed host-reservoir. Correspondingly, the observed coupled variations between $^{3}He/^{4}He$ ratios and $\mu^{182}W$(6, 8) would simply reflect variable proportions of such a hybrid lower mantle reservoir in ascending mantle plumes. In this regard, our model provides an alternative explanation for the origin of negative ^{182}W isotope anomalies in modern OIBs and bridges ^{182}W isotope systematics in Archean mantle derived rocks with observations from modern-day mantle plumes.

Our discovery of long-term preservation of Hadean protocrust in Earth's mantle has also other far-reaching implications, in that their presence requires silicate reservoirs on Earth to have already differentiated during the lifetime of ^{182}Hf(11). Hence, this finding is in support of previous studies(52, 53) that argued the Earth-Moon system to have formed within the first ~60 Ma of our solar system, much earlier than previously thought.

Materials and Methods

Lower Hadean protocrust delamination model

The starting compositions, applied partition coefficients and respective mineral assemblages are listed in SI Appendix, Table S3. In our model we always used internally consistent sets of partition coefficients and assumed batch melting throughout(54). Isotope compositions for ^{143}Nd and ^{176}Hf were modeled by using parent-daughter ratios from the calculated sources, the appropriate decay constants(55–57) and assuming CHUR composition for the BSE(58). The isotope compositions for ^{142}Nd and ^{182}W were back calculated by using the appropriate decay constants(59, 60), present-day isotope composition for the BSE(61, 62), elemental Hf/W and Sm/Nd ratios for the BSE(21, 63), and solar-system initials for the parent-daughter ratios(64, 65). Formation of a mafic protocrust is stage 1 of our model (Fig 3a). The older boundary for the timing of protocrust formation is set by core formation, which could have been completed as early as 38 Ma After solar system formation(21). We assume extraction of mafic protocrust 50 Ma after solar system formation from a mantle with BSE composition(63). The timing of protocrust formation particularly affects the isotope compositions of the short-lived isotope systems during further protocrust evolution (Fig. 5a). For protocrust formation, we used a consistent set of experimental partition coefficients for REE, HFSE, and Th assuming 20% batch

melting at 2 GPa(36). Partition coefficients for W are often incomplete in the literature. If not available, we calculated partition coefficients for W by using partition coefficients for mineral phases from experiments on garnet lherzolite(66) that were adjusted to the melt conditions in our model by using appropriate partition coefficients for Th. Both elements were shown to behave similarly incompatible during silicate crystal-liquid fractionation(21). Hence, the W/Th ratio of our modeled melts extracted from the primitive mantle (W/Th = 0.14) is indistinguishable from the canonical range reported in the literature(21). During stage 2 (Fig. 4b) we re-melt our modeled mafic protocrust and calculate (based on an experimental study(43)) the composition of a typical garnet-rich restite that remained after lower crustal anatexis of a metamorphosed basaltic assemblage (estimated to be representative for the Hadean protocrust) at 12 kbar, in equilibrium with ca. 21% tonalitic melt. The timing of TTG formation as well as the residual mineral assemblage exerts a strong influence on the ^{142}Nd evolution. In contrast, the ^{182}W isotope composition will not change because the ^{182}Hf-^{182}W system went functionally extinct shortly after protocrust formation at ca. 60 Ma after solar system formation. Due to the enriched composition of the precursor and the long half lifes of their parent isotopes, prolonged tonalite formation will only cause small variations in isotopic ingrowth for ^{143}Nd and ^{176}Hf in the lower crustal restites (Fig. 5a). Therefore, prolonged tonalite formation can explain decoupling of ^{142}Nd from the other isotope systems in the residual garnet-rich restites and provides an explanation why ^{143}Nd and ^{176}Hf correlate so tightly with ^{182}W but not ^{142}Nd. Indeed, Hf isotope data in zircon reported for Paleoarchean grey gneisses of the eastern Kaapvaal Craton reveal incorporation of older continental crustal rocks with Eoarchean to late Hadean age(44–46). Assuming prolonged tonalite formation initiated by ca. 4.35 Ga and continued for 100 Myrs the variation of ^{142}Nd isotope composition within their restites would be ca. 9 μ units at 3.55 Ga. In contrast, ^{143}Nd and ^{176}Hf would not vary by more than 1 ε unit.

It has previously been shown that delamination of crustal restites into depleted mantle can cause melting of hybrid mantle sources and that the resulting trace element signatures resemble those of typical komatiites(41). Likewise, the mechanical incorporation of 10-20 % of the garnet-rich restite into an ambient depleted mantle at ca. 3.55 Ga and subsequent 20-30% batch melting of the hybrid source can reproduce the trace element compositions of the SGR komatiites (SI Appendix, Fig. S6). We attribute the variation within the SGR komatiite suite and their partially more depleted trace element compositions, compared to the modeled patterns, to olivine accumulation as indicated by co-variations between MgO content and incompatible trace element concentrations (SI Appendix, Fig. S3). For the ambient depleted mantle into which the restites were mixed we assume 10% melt depletion at ca. 3.85 Ga, using coherent, experimentally constrained melting parameters(36) to calculate the isotope and trace element composition at 3.55 Ga, the age of our samples. The modeled ambient mantle displays initial ε^{143}Nd and ε^{176}Hf values of +4.3 and +7.9, respectively, which is in perfect agreement with the modern terrestrial Hf-Nd array(33), within the range of the modeled DMM composition at 3.55 Ga(67) and also consistent with observational constraints from mantle derived rocks from the Kaapvaal Craton(39). The isotope compositions for ^{182}W, ^{142}Nd, ^{143}Nd, and ^{176}Hf of 20-30% melts extracted from a hybrid mantle source, containing 10-20% of restite and corresponding proportions of ambient depleted mantle, is in accord with the range of isotope compositions observed in SGR komatiites (see Fig. 5b). It is noteworthy that the ^{182}W

isotope composition of the melt is controlled by the restite because high modal abundances of garnet and amphibole, together with refractory Ti-rich phases (rutile/ilmenite), result in high bulk partition coefficients for W. This buffers the ^{182}W isotope composition against possible variations in the ambient Archean mantle (on average ca. +13 ppm(14, 20)). Variable proportions of rutile or ilmenite as a residual Ti-rich phase in the restites do not significantly affect the results of our model. Ratios of Nb/Ta have been proven to be valuable indicators to discriminate between rutile and ilmenite(68), but unfortunately, no Ta concentrations are available for SGR komatiites. We expect ilmenite being present in the restites as this results in reasonable Nb/Th ratios (Nb/Th = 14.8-15.1) that are similar to the range observed in the SGR komatiites (Nb/Th = 11.2-14.6). Evidence for the presence of lower crustal restites in the mantle source of the SGR komatiites is also provided by Zr/Sm (and Hf/Sm) ratios that are best explained by fractionation of garnet. The co-variation of Zr/Sm (and Hf/Sm) with ^{182}W isotope composition is perfectly reproduced by our model (Fig. 2b).

The hybrid mantle source model can also explain the ^{187}Re-^{187}Os isotope inventory of the SGR komatiites. The positive initial $\gamma^{187}Os$ of the SGR komatiites ($\gamma^{187}Os$ = +3.7 ± 0.3) require that their mantle source evolved with a time-integrated suprachondritic Re/Os(25). In our model, the early formation of an enriched protocrust (at ca. 4.52 Ga) with elevated Re/Os ratios caused strongly radiogenic ingrowth of ^{187}Os for ca. 150-250 Ma. As a result, the garnet-rich restites had inherited strongly elevated initial $\gamma^{187}Os$ values at the time of TTG extraction (4.35-4.25 Ga). For Re/Os, the depletion of the restite during TTG melt formation is buffered by the high modal abundance of garnet for which Re was previously shown to behave highly compatible(69, 70). Therefore, the radiogenic ^{187}Os isotope composition will be continuously retained by relatively high Re/Os ratios in the garnet-rich restite. Moreover, overall PGE abundances in the restite should be low as sulfides are expected have melted out during TTG extraction, thus explaining the low PGE abundances in SGR komatiites(25). These conceptual assumptions are also in good agreement with previous estimates that calculated the effects of incorporation of Hadean crust or restites into the source of the SGR komatiites on its Os isotopic composition(25, 71).

It is difficult to assess if Hadean restites that were recycled into the mantle carry a diagnostic Pb isotope composition. Based on mineral partition coefficients alone it is expected that restites, which remained after partial melting of a hydrated mafic protocrust, would develop towards an unradiogenic $^{206}Pb/^{204}Pb$ composition. However, Pb isotope systematics during partial melting of hydrated oceanic crust are not only entirely controlled by mineral partition coefficients as the minerals involved (U, Th and Pb) display dissimilar redox sensitivities and reveal a different mobility in the presence of fluids(72). It is therefore very likely that U-Pb isotope systematics were often affected by metasomatism resulting in highly variable initial Pb isotope compositions in Archean TTGs and their mafic counterparts. Indeed, while many Archean cratons have preserved a long-lived high-μ continental lithosphere with distinctive Pb-isotope compositions(73), other cratons show large Pb isotope variations with more unradiogenic Pb isotope patterns(74–76).

Assessment of magma ocean models involving perovskite fractionation

Alternative models have been proposed for the origin of the SGR komatiites(12), involving fractionation of a high-pressure and temperature Mg- and Ca-perovskite mineral assemblage in an early terrestrial magma ocean. In short, we found that such models are highly dependent on the sets of partition coefficients used. Independent of this issue, some important diagnostic features of the samples analyzed here (e.g., ^{176}Hf-^{143}Nd-^{138}Ce isotope relationships) cannot be reproduced by a magma ocean model or do not require the presence of perovskite cumulates.

In detail, we modeled evolution of a mantle reservoir that has undergone perovskite segregation, tightly following a previous model for the SGR(12). Herein, a primitive mantle undergoes removal of 10% Mg- and Ca-perovskite at 4.537 Ga before it evolves until 4.027 Ga. Subsequently, this reservoir undergoes batch melting in the spinel stability field at 4.027 Ga before it melts at 3.55 Ga to produce the SGR komatiites. For perovskite, we used a more rigorous, internally consistent set of partition coefficients from reference 35 using laser ICPMS data for their representative experiment H2020 a+b and relative abundances of Mg and Ca perovskite of reference 12. Mantle depletion at 4.027 Ga was modeled in analogy to the parameters presented by reference 12, but again using more updated sets of partition coefficients(36). All other parameters like decay constants or CHUR values are as above. We refrained from modeling W because the original dataset for perovskite(35) does not include W partition coefficients. Previous modeling(12) referred to lattice strain modeling of D_W, but the lattice strain model used by reference 35 is only applicable to 1+, 2+, 3+, and 4+ ions. Recent work(77, 78) has shown that the valence state of W, even in the more reduced regime of an early magma ocean, is rather 6+.

Hafnium-Nd-Ce modeling results are shown as blue symbols in Fig. 3, recalculated to 3.55 Ga. Notably, none of the modeled composition matches compositions of the SGR komatiites in Hf-Nd and Nd-Ce space. It is important to note, that the fractionation of perovskite during magma ocean crystallization does not lead to suprachondritic Lu/Hf and to a decoupling of ε^{143}Nd and ε^{176}Hf systematics, once recalculated to 3.55 Ga. Rather, the Hf-Nd composition of the modeled mantle after perovskite segregation is near chondritic at 3.55 Ga. Depletion of such a mantle reservoir at 4.027 Ga in the spinel stability field (“shallow mantle”) yields decoupled Hf-Nd isotope compositions, but at extremely radiogenic ε^{143}Nd at a given ε^{176}Hf, which is nowhere found in our sample set. Larger amounts of residual garnet during mantle depletion at 4.027 Ga (“deep mantle”) may result in ε^{143}Nd-ε^{176}Hf systematics that resemble the compositions of SGR komatiites, in analogy to the modern-day terrestrial mantle array(34). Regardless of this consideration, magma ocean models involving perovskite fractionation are not required to explain combined ^{143}Nd-^{176}Hf systematics of samples from the Kaapvaal Craton. In Nd-Ce isotope space, both mantle reservoirs (“shallow & deep mantle”) yield much less radiogenic ε^{138}Ce at 3.55 Ga than found in our samples.

Collectively, our modeling of a mantle reservoir involving perovskite segregation and subsequent mantle depletion cannot easily explain compositions found in our sample set and previously published isotope data for the SGR.

Analytical protocol

Our analytical protocol for isotope dilution analysis follows procedures that were described in detail by previous studies(20, 53, 79). For ^{138}La-^{138}Ce measurements we processed 1g of sample powder. For La-Ce isotope dilution (ID) measurements a 5% aliquot was spiked with a ^{138}La-^{142}Ce isotope tracer. For the 95% aliquot we utilized the first stage cation resin column of a previously published protocol for W(14) to separate REE from matrix elements for high precision ^{138}Ce isotope composition (IC) measurements. This step is required since sample loads larger than 200mg exceed the capacity of the first stage column in our ^{138}La-^{138}Ce separation protocol(79). Measurement protocols for La-Ce ID measurements as well as for Ce IC measurements followed a previously described routine(79) except that 10^{12} Ω resistors used for interference corrections were replaced by 10^{13} Ω resistors. All data were normalized relative to ^{136}Ce/^{140}Ce of 0.002124072(80) and are given relative to a ^{138}Ce/^{136}Ce value of 1.33738 for the Mainz AMES standard solution(81). All samples were analyzed repeatedly. Reported uncertainties either refer to the corresponding 95% CI (n ≥ 4) or to our intermediate precision (± 0.21 ε-units)(38).

High-precision ^{182}W isotope measurements mainly followed established analytical protocols(14, 20) that were slightly modified to yield highly purified W solutions from large sample loads (up to 18g) and to improve our analytical uncertainty. In short, samples were measured at average signal intensities of 17 V for ^{182}W (using 10^{11} Ohm amplifiers) corresponding to a ~175 ng/ml W sample solution at an uptake rate of ca. 55 µl/min. Samples were always bracketed by a concentration-matched certified reference material (NIST SRM 3163). Results of high precision W isotope analyses are reported in the µ-notation (equivalent to ppm) relative to the bracketing NIST solutions and always refer to the measured ^{182}W/^{184}W ratio that has been corrected for mass bias by using ^{186}W/^{184}W = 0.92767(82). All samples were repeatedly analyzed (n=6-11) and uncertainties for average W isotope compositions are correspondingly reported as 95% confidence intervals (see SI Appendix, Table S1).

Our protocol for the chemical purification of W for high-precision isotope composition analysis comprises four columns. During a cation (AG 50 W-X8 resin, column I) and anion exchange stage (AG 1-X8 resin, column II) W is separated from matrix elements and HFSE & Ti, respectively. Columns III (TEVA resin) and IV (TODGA resin) are clean up columns that yield purified W cuts. In this regard, the repetition of the final stage column during the chemical separation of W(20) improves the purification from remaining matrix elements. The final W-bearing eluate was directly loaded onto BioRad Poly-Prep® columns filled with 0.8ml Eichrom prefilter® material to extract organic compounds. This, together with threefold treatments with 80 µl of cHNO_3-30%H_2O_2 at max. 60°C after dry-down steps during and after the chemical separation, strongly improved yields and removed mass independent effects on ^{183}W(14). Prior to loading onto our final stage column, we combined up to 10 cuts in case sample powders were split up into aliquots (up to 1.3g) during matrix separation. The combination of sample solutions during chemical separation does not affect the accuracy of our high precision ^{182}W isotope analysis as demonstrated by indistinguishable results for sample solutions of our in-house rock reference material LP 1 (historical La Palma Basalt), that were either obtained from single column cuts (up to 1.3g) or combined solutions from 10 column cuts (in total 11.3 g). The purpose of combining the final cuts is to efficiently measure the cuts by reducing the cumulative

volume of leftovers after multiple measurements of individual solutions. This allows measuring at the highest beam intensities possible and, together with our refined separation procedure, significantly improves the analytical uncertainty of our measurements. This is also reflected by our intermediate precision of our in-house rock reference materials LP 1 and AGC 351 that were always measured in every session, yielding markedly improved 2 SD of ± 1.5 ppm and ± 2.7 ppm, respectively (SI Appendix, Fig. S1). The $\mu^{182}W$ session averages for LP 1 (1480 OIB from La Palma) and AGC 351 (3455 Ma gneiss from Swaziland) overlap within their 95% CI (LP1 = -0.4 ± 1.0 ppm and AGC 351 = -0.2 ± 0.5 ppm) and are indistinguishable from the NIST reference material and previously reported long-term averages for the same sample powders(14, 20). Additionally, we also performed repeated analyses (n = 15) of a 3.27 Ga old Komatiite (sample 160245, Ruth Well Formation) from the Pilbara Craton Western Australia that exhibits highly elevated W concentrations of 19.1 µg/g(20). The $\mu^{182}W$ session average for sample 160245 ($\mu^{182}W$ = +7.9 ± 0.7 ppm, 95% CI) is in agreement with previous results(20) and shows a good intermediate precision (2 SD of ±2.5 ppm). This, together with the elevated ^{182}W isotope composition and high W concentration of sample 160245 validates the method for analytical campaigns addressing ^{182}W isotope systematics in Archean mantle derived rocks that often display anomalous ^{182}W isotope compositions.

Competing interest statement

The authors declare no competing interest.

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Figures

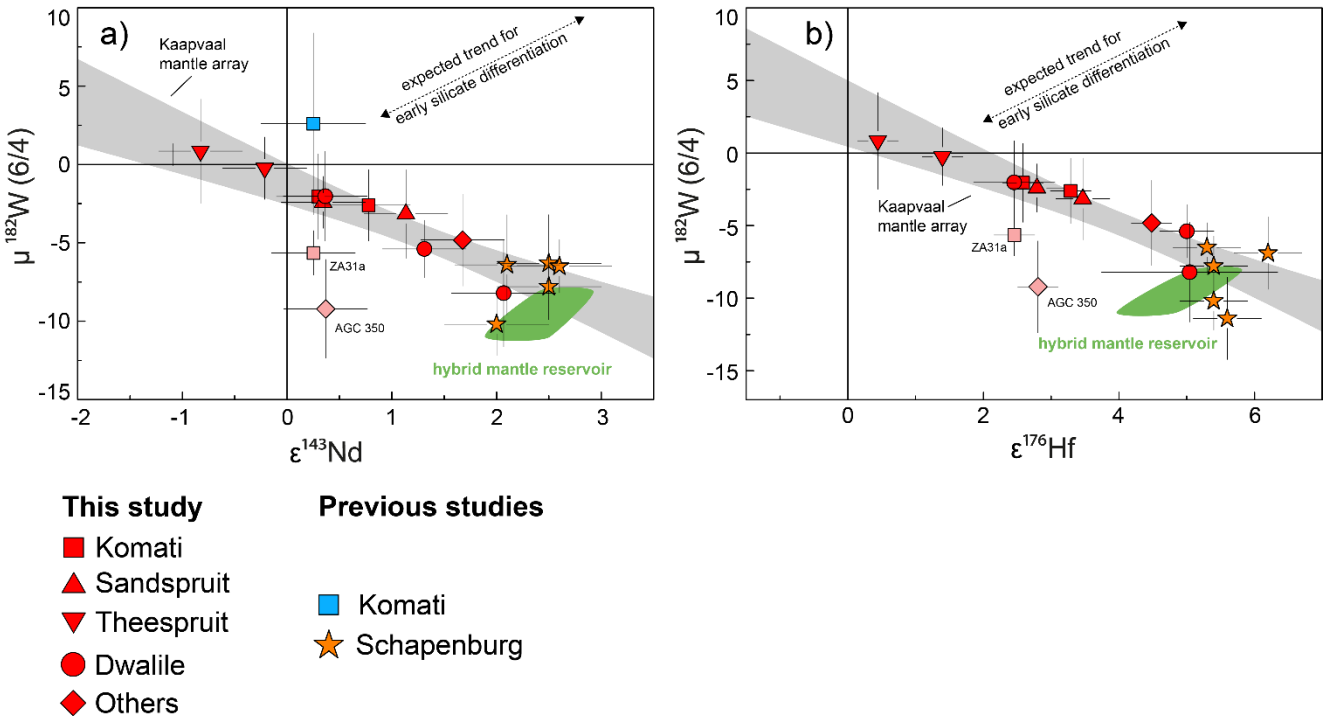


Fig. 1: Measured $\mu^{182}W$ vs. $\varepsilon^{143}Nd_{(t)}$ (a) and $\mu^{182}W$ vs. $\varepsilon^{176}Hf_{(t)}$ (b) for mantle derived mafic rock samples from the Kaapvaal Craton including literature data. The ^{182}W isotope composition for sample AGC 350 and ZA31a (pale red symbols) were most likely overprinted by metasomatic agents carrying negative ^{182}W isotope compositions. The literature data include previously published data for komatiites from the Schapenburg Greenstone Remnant (orange asterisks)(12) and the Komati Formation (blue square)(11, 16). We note that previously published literature data for the Komati Formation only report combined $\mu^{182}W$ vs. $\varepsilon^{143}Nd_{(t)}$ data for one single sample (sample BV 02, blue square)(11, 16). The green fields illustrate modeled values of our proposed hybrid reservoir (10-20% restites admixed to ambient depleted mantle). The shaded grey field, referred to as *Kaapvaal mantle array*, is an uncertainty envelope employing the 95% confidence interval in which of all mantle derived samples are expected to fall. Note, that the negative co-variation displayed by the *Kaapvaal mantle array* does not follow the expected trend for early silicate differentation (indicated by dashed line in panel 1b).

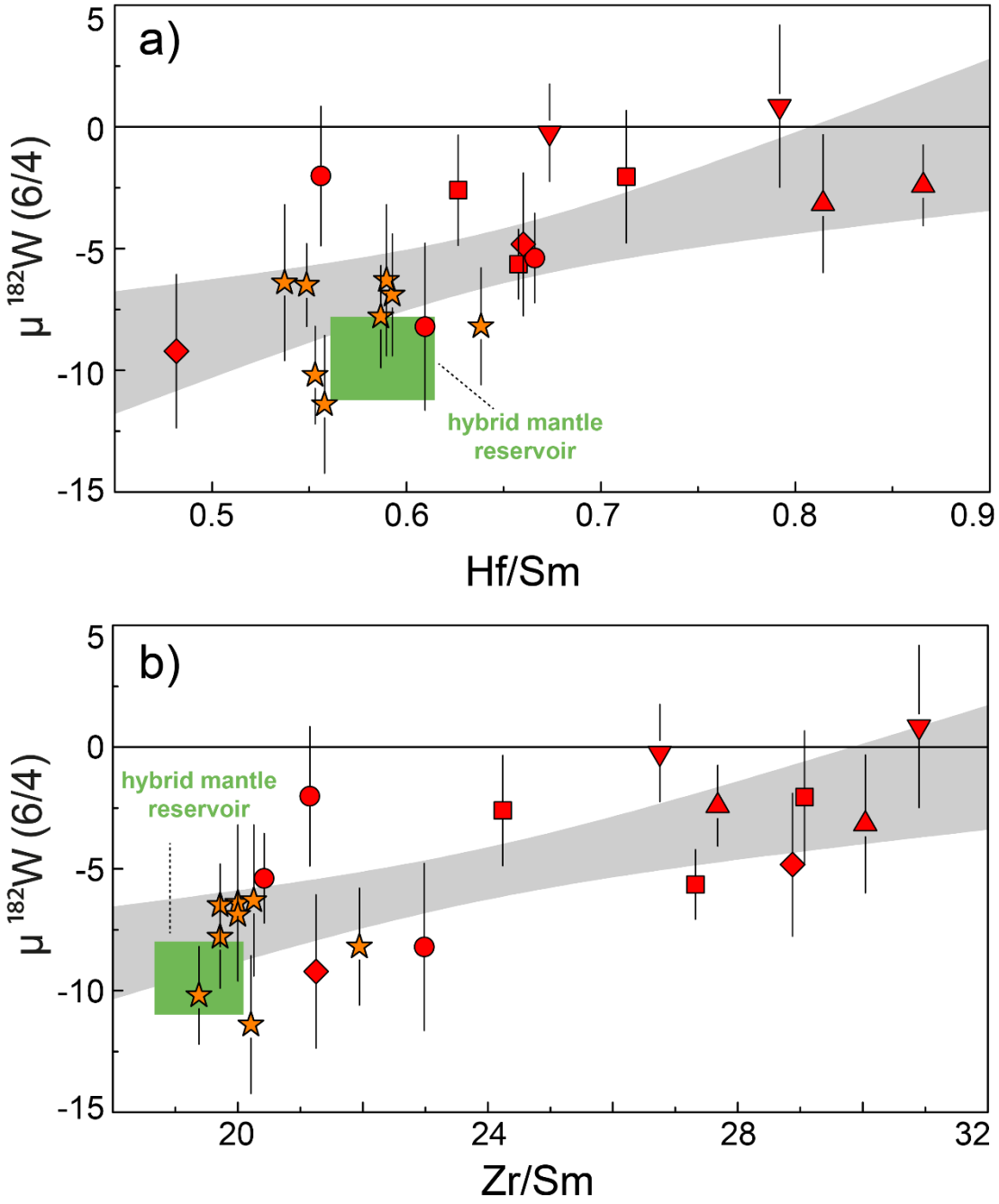


Fig. 2: Plot of μ^{182}W vs. (a) Hf/Sm and (b) Zr/Sm for rocks from the Kaapvaal craton. Symbols are the same as in Fig. 1. Data for komatiites from the Schapenburg Greenstone Remnant (SGR) are taken from the literature(12). The combined data indicate a systematic co-variation between ^{182}W isotope composition and Hf/Sm, Zr/Sm ratios with one endmember defined by the SGR komatiites. The negative μ^{182}W anomalies and low Hf/Sm and low Zr/Sm ratios prominent in the SGR komatiites can be attributed to the presence of 10 – 20% garnet rich restites within a hybrid source that underwent 20 – 30% batch melting (green box). The grey shaded array refers to the 95% confidence intervall in which of all samples are expected to fall.

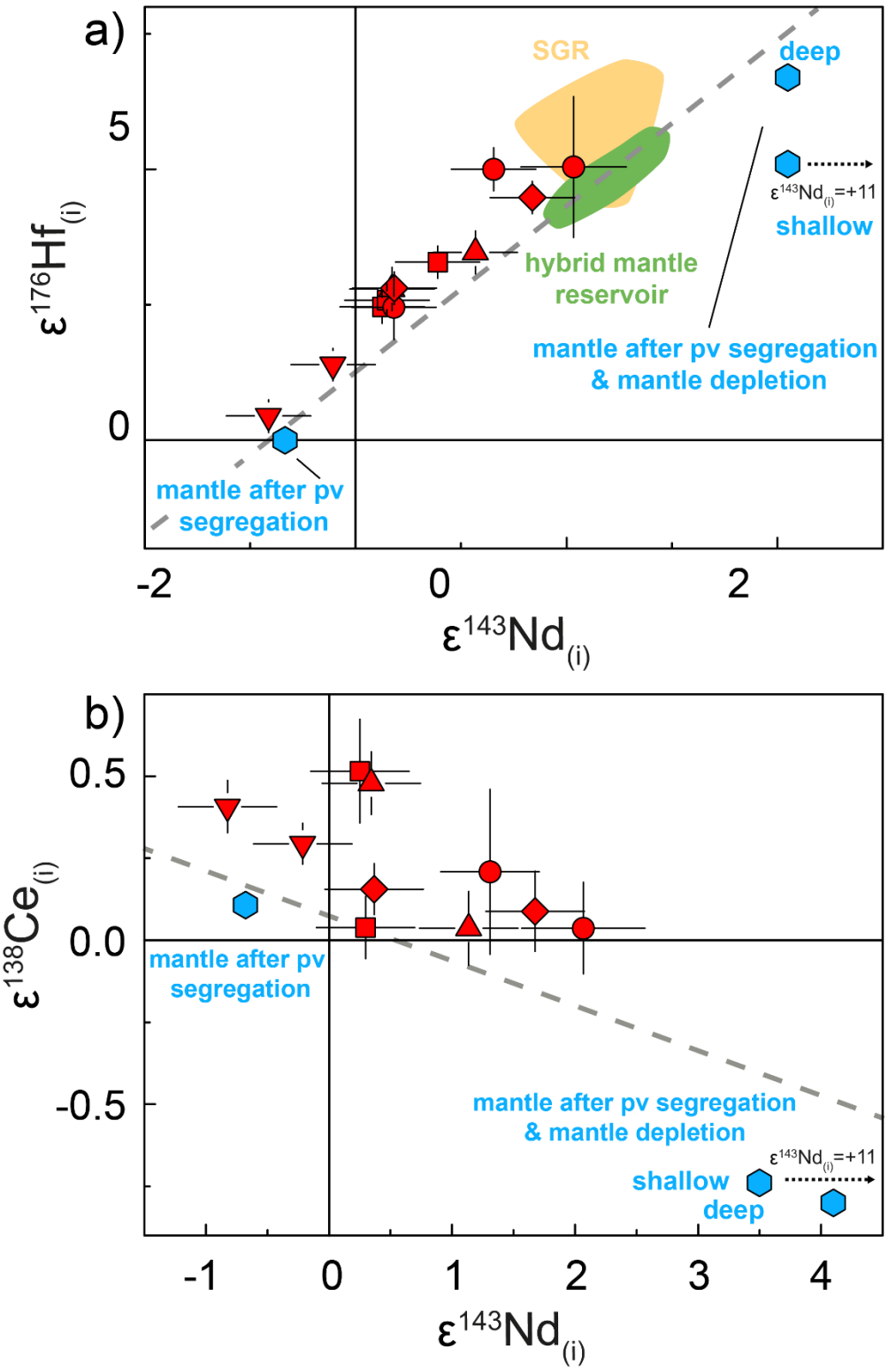
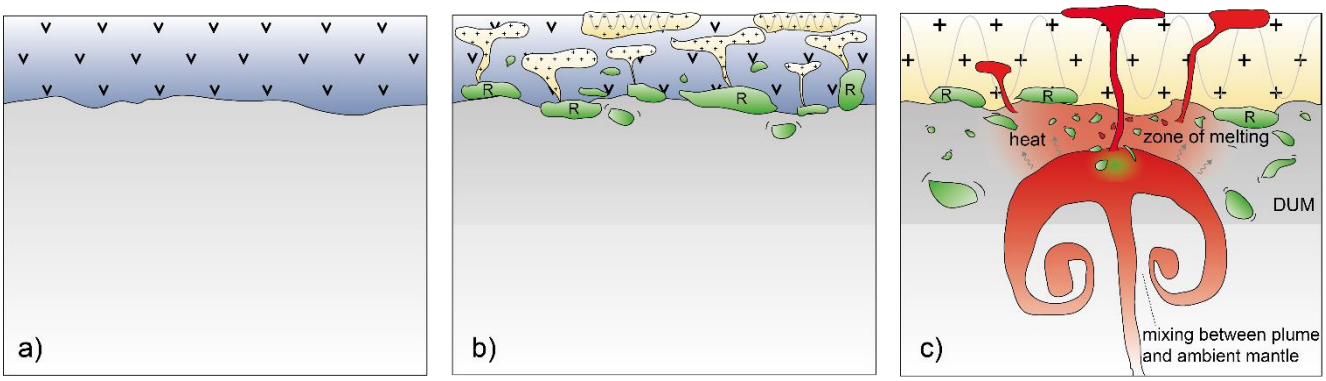


Fig. 3: Plot of $\varepsilon^{176}Hf_{(t)}$ vs. $\varepsilon^{143}Nd_{(t)}$ (a) and $\varepsilon^{143}Nd_{(t)}$ vs. $\varepsilon^{138}Ce_{(t)}$ (b) for mantle-derived mafic-ultramafic rocks from the Kaapvaal Craton analyzed in this study (red symbols). Symbols are the same as in Fig. 1. The orange field marks previously published compositions of komatiites from the Schapenburg Greenstone Remnant (SGR)(12). The green fields illustrate modelled values of our proposed hybrid mantle reservoir (10-20% restites admixed to ambient depleted mantle). Blue symbols illustrate our modelling results for mantle reservoirs (at 3.55 Ga) that underwent perovskite segregation and subsequent melt depletion (using an internally consistent set of partition coefficient, see method section), either in the spinel stability field ("shallow") or at the presence of garnet ("deep"), illustrating that there is no clear evidence for involving reservoirs that underwent perovskite fractionation in an early magma ocean. The dashed line shows the modern Terrestrial Array for MORBs and OIBs ($\varepsilon Hf = 1.55 \times \varepsilon Nd + 1.21$ and $\varepsilon Ce = -0.14 \times \varepsilon Nd + 0.05$)(33, 37). The $\varepsilon^{138}Ce_{(t)}$ of the modeled hybrid reservoir is not shown in (b) due to large modeling uncertainties imparted by the highly incompatible behavior of La and Ce.



Stage 1 - Protocrust Formation
4.517 Ga

Stage 2 - TTG & Restite Formation
4.35 - 4.25 Ga

Stage 3 - hybrid source melting
3.55 Ga

Fig. 4: Preferred geodynamic model for the origin of crustal and mantle-derived rocks from the Kaapvaal Craton. **(a)** Formation of a mafic protocrust by ca. 50 Ma after solar system formation. **(b)** Formation of TTG-like batholiths (orange) and residual garnet – rich restites (green, labelled „*R*“) after partial protocrustal anatexis between ca. 4.35 and 4.25 Ga. **(c)** Recycling of lower crustal restites and plume initiated volcanism lead to melting of hybrid sources that involved delaminated restites, depleted ambient upper mantle and primitive mantle supplied by the ascending plume. Shades of grey visualize depleted upper mantle (DUM) and lower mantle (light grey).

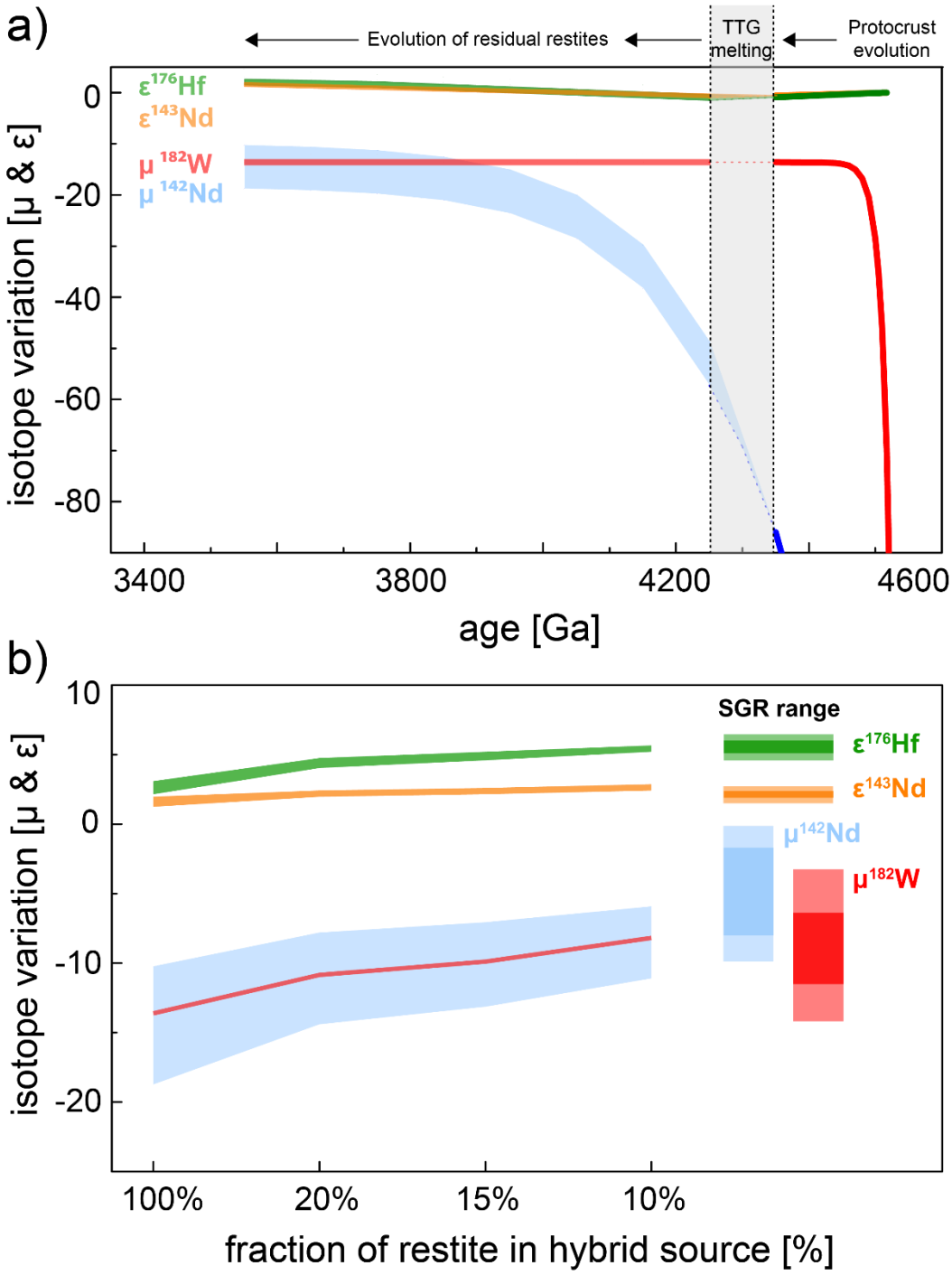


Fig. 5: Isotope evolution graphs for the proposed geodynamic model involving mantle recycling of lower crustal restites. (a) During stage 1, mafic protocrust formed ca. 50 Ma after solar system formation developed strongly unradiogenic isotope compositions, in particular for ^{182}W and ^{142}Nd. Stage 2 marks restite formation during TTG extraction from mafic protocrust. The grey bar ilustrates the time interval (4.35 to 4.25 Ga) over which TTG extraction affects the isotope compositions of residual restites. Depending on the exact timing of TTG extraction, the restites develop to markedly different ^{142}Nd isotope composition with time (blue field). In contrast, ^{182}W is insensitive to the timing of TTG extraction, because ^{182}Hf went extinct shortly after formation of the mafic protocrust. Due to their longer half lifes the effects on long-lived radionuclides are rather negligible. **(b)** Mixing calculations illustrating the isotope composition of the proposed hybrid reservoirs as a function of delaminated restites mixed into ambient depleted mantle. Ca. 10-20% of admixed restite to ambient depleted mantle reproduces the isotope compositions found in the SGR endmember. This hybrid source mixed with primitive mantle material supplied by ascending mantle plumes as reflected in the *Kaapvaal mantle array* for ^{182}W and long-lived radiogenic nuclides (see Fig. 1).

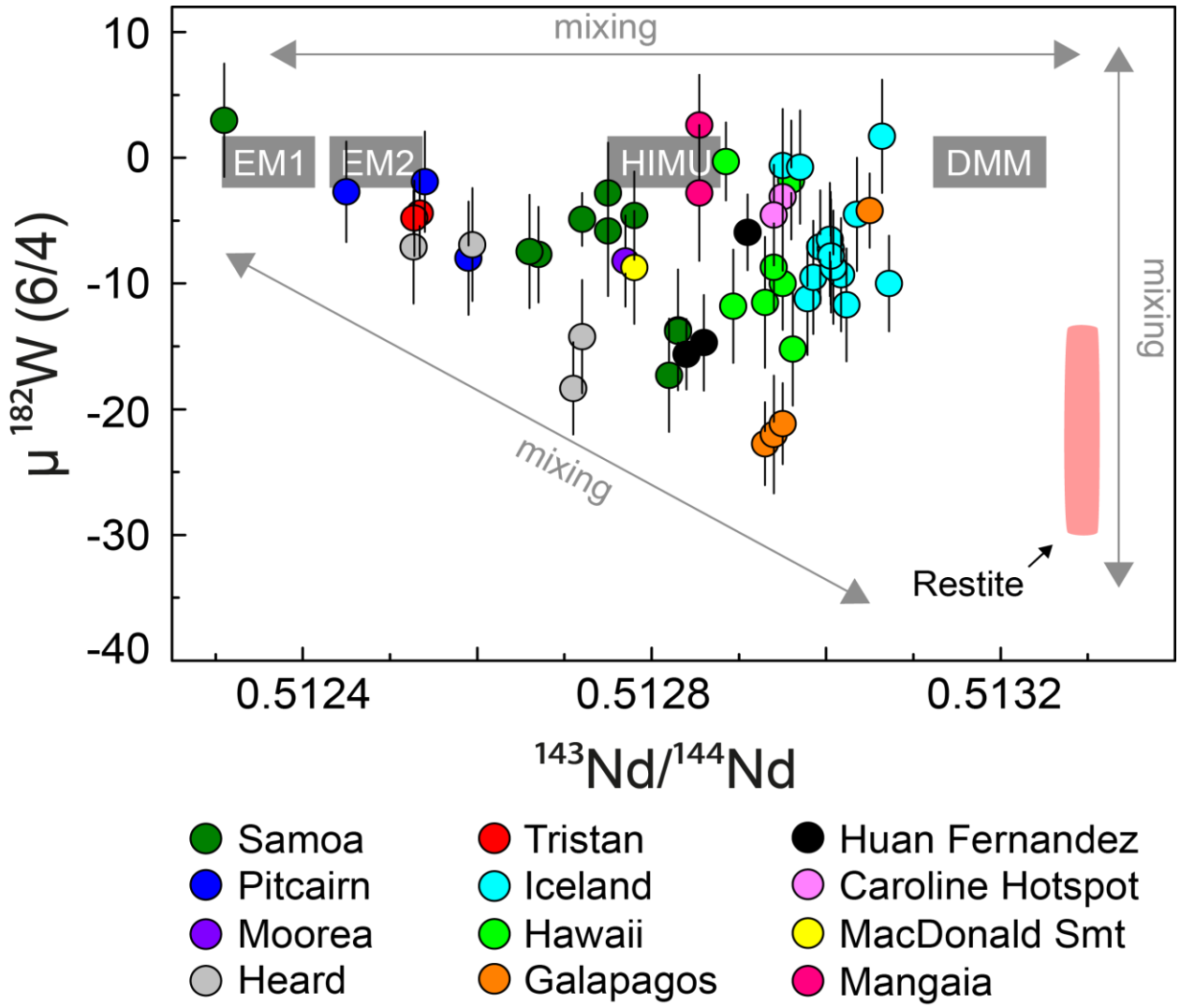


Fig. 6: Compilation of combined ^{182}W and ^{143}Nd isotope data available for modern OIBs. Data were compiled from recent studies(8, 48) and references therein. Notably, the global compilation for modern OIBs shows a similar pattern than the Archean mantle-derived rock assemblage from the Kaapvaal Craton including an endmember with low μ^{182}W and radiogenic ^{143}Nd/^{144}Nd. Also shown is the present ^{182}W and ^{143}Nd isotope composition calculated for Hadean restites that remained isolated in the mantle (pink field). This global dataset can be best explained by the admixture of the classical mantle endmember components DMM (depleted MORB mantle), EM1 (enriched mantle I), EM2 (enriched mantle II) and HIMU (high "μ" or high ^{238}U/^{204}Pb) to a primordial reservoir that is characterized by negative ^{182}W anomalies and depleted ^{143}Nd isotope composition. A similar OIB compilation for ^{182}W and ^{176}Hf is limited by the availability of ^{176}Hf isotope data but shown for comparison in SI Appendix, Fig. S8.

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SI Appendix

Geological background of our sample selection

We analyzed a comprehensive set of rocks from the Kaapvaal Craton that range from different types of grey orthogneisses (TTGs and more evolved granitoids) to mantle-derived lithologies of mafic-ultramafic composition. This representative suite of 17 samples span an age range from 3.55 to 3.22 Ga and represent the main lithological units of the Ancient Gneiss Complex (AGC), also comprising the oldest mafic rocks (lower Onverwacht Group, 3.55 to 3.45 Ga) of the Barberton Granite-Greenstone Terrane (BGGT).

The AGC is located in Swaziland and is a typical high-grade gneiss terrain that comprises 3.66-3.20 Ga old rocks(83). The oldest part of the AGC are polydeformed granitoid gneisses, heterogeneous in age and composition(46, 84), that are interbanded with amphibolites. Together, they formed layered grey gneiss sequences in response to ductile deformation under high strain conditions(83). The different varieties of rocks from this sequence have been summarized as the Ngwane Gneiss (NG)(85). The oldest generation of NG (NG *sensu stricto)* are 3.66 Ga to 3.5 Ga granitoid gneisses(44–46, 84, 86) that mainly belong to the tonalite-trondhjemite-granodiorite (TTG) suite but also comprise granitic rocks. As indicated by trace element systematics(44), whole rock Nd isotope systematics(87) and Hf-in-zircon isotope data(44–46) the protoliths of the orthogneisses resulted, at least in part, from melting of a LREE enriched source with considerable residence time, most likely older continental crust of Eoarchean to late Hadean age. Younger generations of grey gneisses, which are mapped as NG, were emplaced after 3.45 Ga. These show the same field appearance as the 3.66-3.45 Ga NG but are as young as 3.2 Ga(87, 88). The oldest NG hosts scattered remnants of supracrustal assemblages with greenstone belts (e.g. Dwalile Supracrustal Suite, DSS(89)). These remnants postdate the oldest NG, vary in size and are either infolded, occur as tectonically intercalated xenoliths of a few centimeters or even represent coherent blocks of several kilometers(89, 90). The origin of these remnants remains contentious. They were interpreted either as strongly flattened dikes(89, 91) or as dismembered portions of the Dwalile Greenstone Remnant (DGR), which represents the largest of the greenstone remnants of the AGC(46, 83, 90, 92). The DGR is located in SW Swaziland and the supracrustal rock assemblage (metavolcanics, metasediments) were shown to be extruded between 3.44 and 3.46 Ga, therefore postdating the oldest generation of NG(19, 90, 93). Notably, the metavolcanic rocks from the DGR share geochemical similarities with volcanic assemblages from the Onverwacht Group which hints at a genetic link between the DGR and the BGGT(19, 90, 91). Based on trace element systematics and variable whole-rock initial εNd and εHf values it has been argued that the mafic and ultramafic DGR rocks were derived from a mildly depleted mantle source and were in part contaminated by rocks from an ancient continental source, presumably crustal material of NG-like composition(19, 90). The oldest NG and intercalated members of the DSS were intruded by the texturally and compositionally distinct Tsawela Gneiss between 3.48-3.43 Ga(45, 88, 89, 94, 95) and younger generations of grey gneisses that date back to ca. 3.2 Ga(88).

All sample localities are shown in SI Appendix, Fig. S7 and GPS coordinates are provided in previous studies(18, 46, 92, 94). We have analyzed two grey gneisses from

the >3.45 Ga NG suite that were collected along the Mtimane River in the Mankayane area in central Swaziland, where granitoid gneisses of different ages were variably affected by intensive regional migmatization at ca. 3.2 Ga(90, 96, 97). Both samples (AGC 351 and AGC 352) were previously described(46, 94). AGC 351 is a 3.455 Ga old, strongly migmatized grey gneiss of near granitic composition and interpreted to be derived from felsic crustal precursors that mixed with juvenile, depleted-mantle derived melts(46, 96). AGC 352 is a 3.442 Ga very homogeneous fine grained grey gneiss(94).

We have analyzed several samples from greenstone remnants that are interlayered with grey gneisses of the AGC. We investigated two komatiites and one amphibolite from the DGR (AGC 83, AGC 86 and AGC 38), one typical amphibolite fragment as found in the AGC (AGC 222) and a 3.455 Ga gabbroic enclave (AGC 350) from central Swaziland. The mafic-ultramafic rock samples from the DGR were previously characterized(19, 90). Sample AGC 222 is a fragmented amphibolite enclave from Kubuta in central Swaziland with a minimum age of 3.4 Ga(98). It is similar in composition to other greenstone remnants found in the AGC(46, 94). Gabbroic enclaves like AGC 350 can be found along the Mtimane River in the Mankayane area close to the sample localities of AGC 351 and AGC 352. As described by reference (96) the precursors of the gabbroic enclaves were emplaced together with granitoid gneisses at 3.455 Ga. At about 3.2 Ga, a tectono-magmatic-metamorphic event reworked the grey gneisses and greenstones(97) which led to boudinage and local anatexis of the gabbros and migmatization of the grey gneisses (e.g. sample locality of AGC 351).

The youngest samples from the AGC are two ca. 3.2 Ga gneisses. Sample AGC 473 is a 3.24 Ga grey gneiss of trondhjemitic composition, which intruded into the oldest generation of NG northwest of the DGR. Based on structural considerations, the adjacent NG were interpreted as basement for the volcanic sequences of the DGR(89). Our younger grey gneiss sample AGC 473 belongs to the youngest generation of NG but contains inherited zircon grains of 3.49 Ga and ca. 3.64 Ga(15). This young generation of grey gneisses belongs to a 3.2 Ga magmatic event that is typically associated with indicators for strong deformation and high-grade metamorphism and therefore suggested to be the result of migmatization and crustal melting of older generations of crustal rocks(92, 96). Sample AGC 445 is a 3.216 Ga old grey gneiss from the Piggs Peak area also belonging to the former 3.2 Ga NG generation(92).

The AGC is in faulted contact with the BGGT along the ca. > 3.2 Ga old Phophonyane shear zone northwest of Pigg´s Peak town(86) and is spatially separated by sheet-like intrusions of the Mpuluzi and Piggs Peak batoliths. Rocks from the BGGT comprise a complex association of greenstone sequences and grey gneisses. The greenstone sequences in the BGGT (referred to as the Barberton Greenstone Belt, BGB) comprise a complex association of volcanic-sedimentary rocks that were deposited over more than 300 million years from < 3547 to > 3219 Ga(99). The volcano-sedimentary sequence of the BGB (known as the Barberton Supergroup) has traditionally been divided (from base to top) into three main lithostratigraphic units: The Onverwacht, Fig Tree, and Moodies groups. The Onverwacht Group (OG) is the oldest greenstone succession of the BGB and comprises voluminous mafic to ultramafic metavolcanics successions with sparsely interbedded metasediments. As we only analyzed samples from the lower OG, we only provide a short overview about the lowest stratigraphy of

the BGB. The OG is subdivided into the lower and upper Onverwacht Group, marked by a chert layer, known as the Middle Marker. The lower OG comprises the Sandspruit, Theespruit, and Komati Formations, the upper OG includes the Hoggenoeg, Noisy, Mendon, and Kromberg Formations(100). The oldest magmatic events preserved in the lithostratigraphic succession of the BGB are mafic-ultramafic and felsic metavolcanic rocks. This bimodal sequence (originally assigned to the Sandspruit and Theespruit Formations) comprises the oldest rocks of the lower Onverwacht Group. The metavolcanic rocks of the Sandspruit and Theespruit Formations were shown to be time-equivalent and deposited during one single volcanic event at ca. 3530 Ma and therefore constitute a single lithostratigraphic unit(101). The record of the somewhat younger 3.482 Ga Komati Formation(100) bears witness to a period of prolonged volcanic activity, as it comprises a continuous succession of alternating komatiitic, komatiitic basalt, and tholeiitic basalt lava flows without any intercalated sedimentary layers that would reflect a hiatus in the stratigraphy(102).

The BGB is surrounded by 3.521 to 3.197 Ga old granitoid gneisses(92) that form a cluster of 12 diapiric plutons with a wide variety of compositional types that intruded into the lowermost formations of the BGB(103). They can be subdivided into two major compositionally families that were emplaced during two periods: The older (3.45-3.2 Ga) TTG group that was coeval with deposition of supracrustal sequences in the BGB, and the much younger (ca. 3.1 Ga) GMS group (granite-monzonite-syenite) which intruded after sedimentation and stabilization of the crust through continued deformation of the TTG basement and greenstone sequences at ca. 3.2 Ga(99).

Our samples were collected at the southwestern margin of the BGGT southeast of the town of Badplaas, in an area around the settlement of Tjakastad (SI Appendix, Fig. S7). Here a significant proportion of the metavolcanic rocks from the Sandspruit and Theespruit Formations occur as dismembered rafts and xenoliths in tonalitic-trondhjemitic gneisses of the Badplaas, Stolzburg and Theespruit Plutons in the southern part of the Barberton Mountain Land(103, 104).

Supplementary Figures

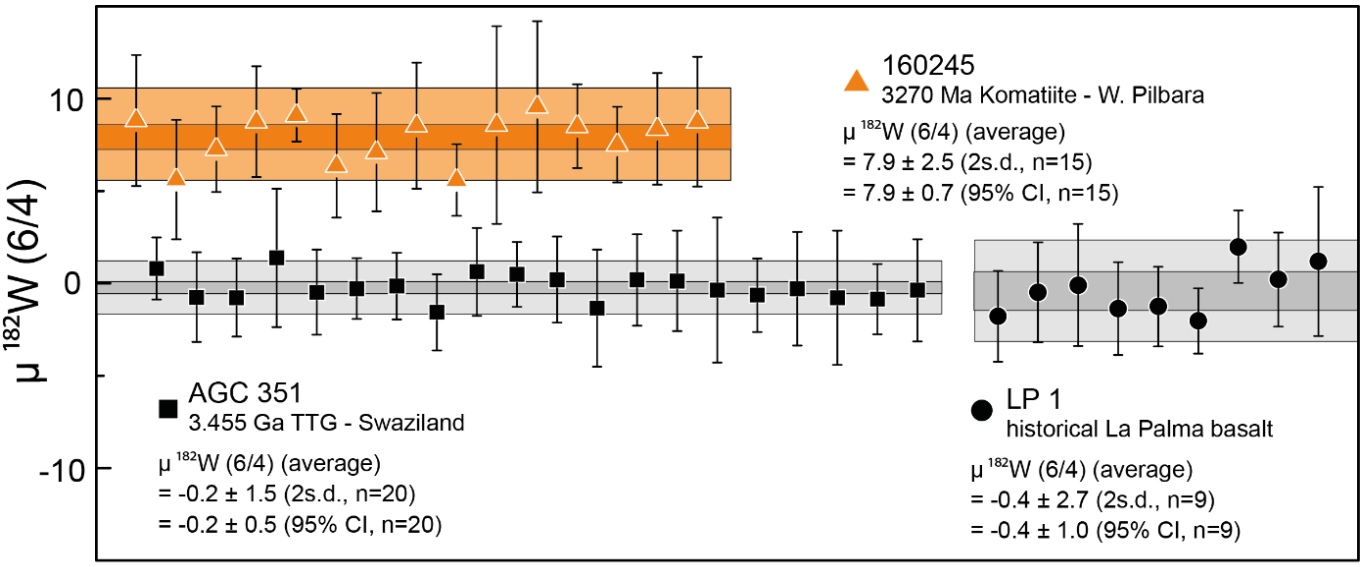

Fig. S1: Intermediate precision for $\mu^{182}W$ (6/4), inferred from the repeated analysis of multiple digestions for our in – house reference materials AGC 351, LP 1, and 160245 that are reported relative to W NIST SRM 3136. Each symbol refers to the average value of multiple measurements conducted during an analytical session. The uncertainties for the session mean values are given by the corresponding 95% CI. The intermediate precision for our in–house reference materials are given by the 2 SD of the session mean values.

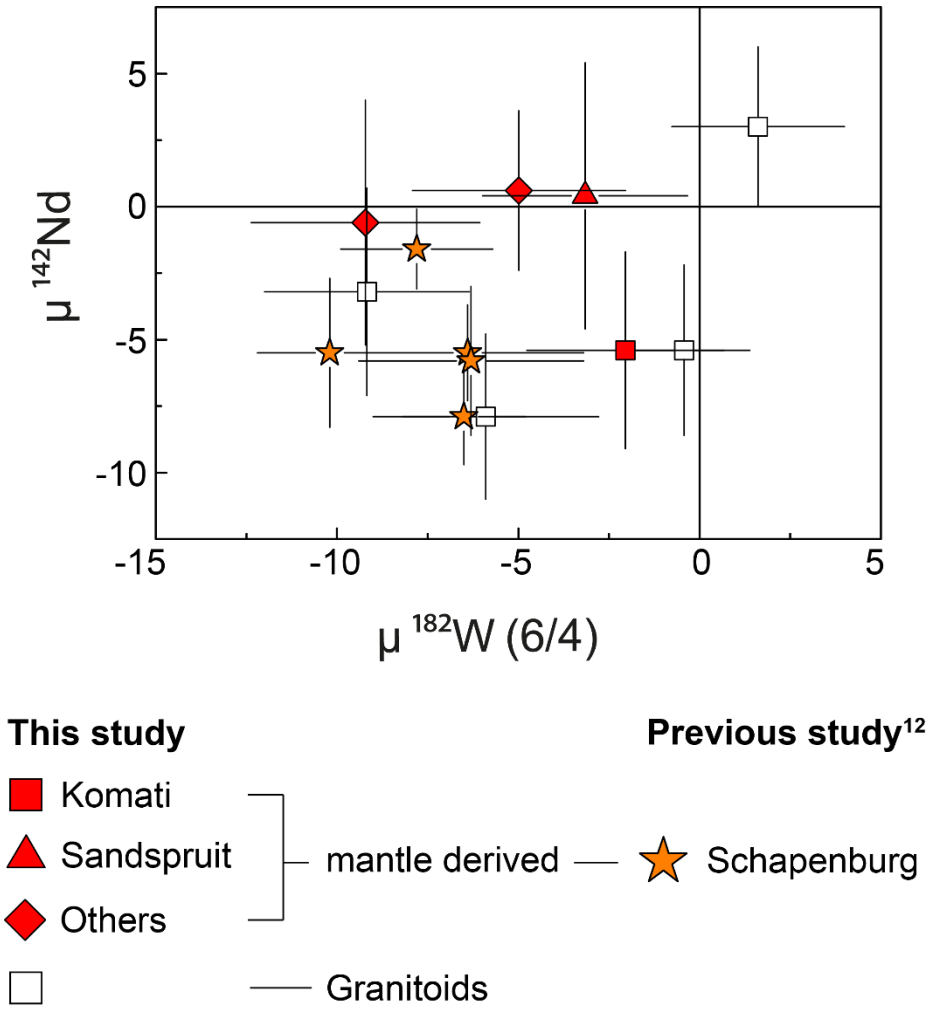

Fig. S2: Compilation of available $\mu^{182}W$ and $\mu^{142}Nd$ for crustal and mantle derived rocks from the Kaapvaal Craton. Symbols for our samples are the same as in Fig. 1. The ^{142}Nd isotope compositions for samples from our study were presiously reported(15) and combined ^{182}W – ^{142}Nd systematics for komatiites from the Schapenburg Greenstone Remnant (SGR) were taken from the literature(12). The combined data show a tendency towards negative anomalies but reveal no clear correlation due to the comparatively large uncertainties.

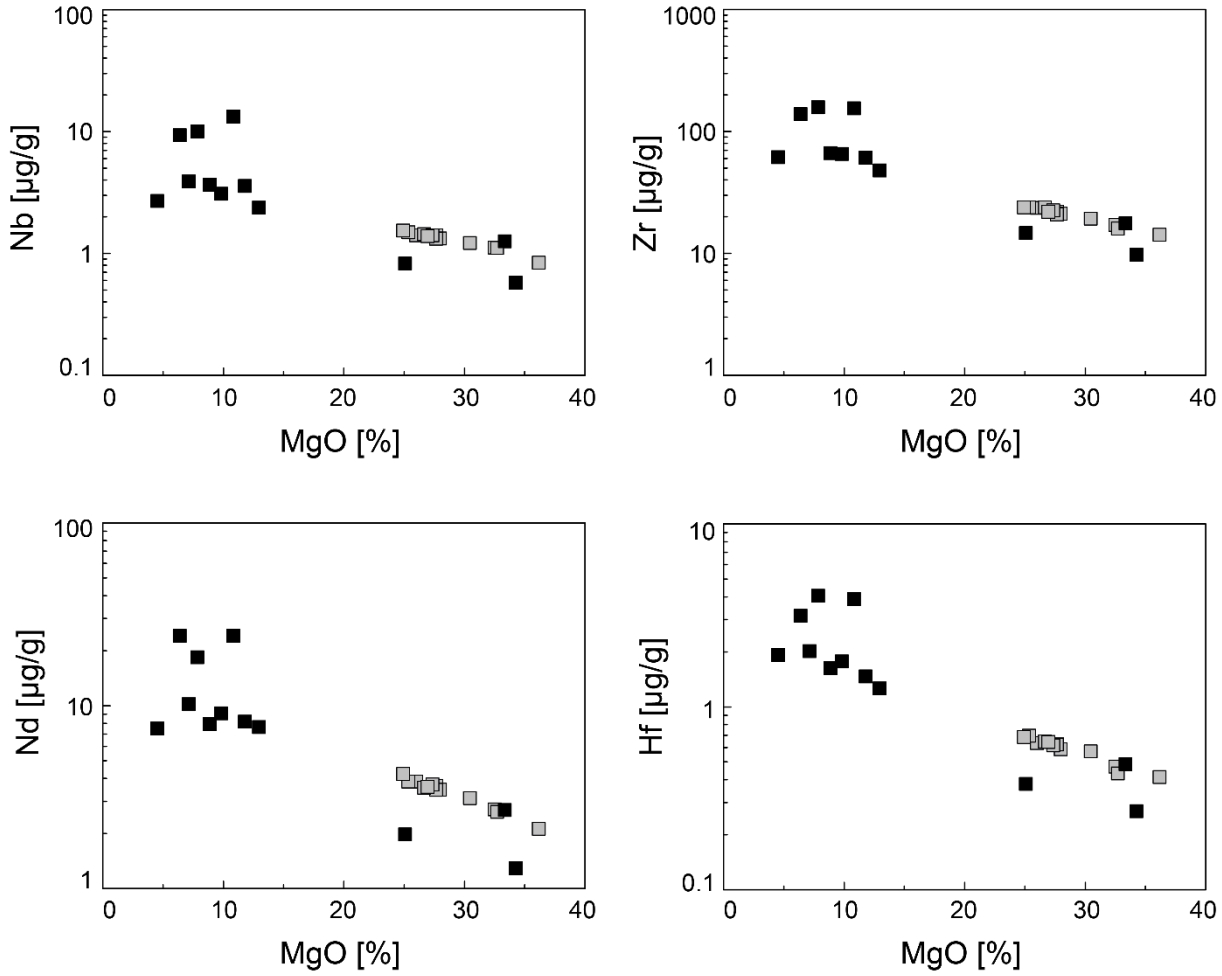

Fig. S3: Trace element variation diagrams [µg/g] vs. MgO-content [%] for our samples and the Schapenburg komatiite suite(12)**.** The inverse correlations of incompatible elements (e.g. Nb, Zr, Nd, Hf) in variation diagrams vs. MgO content reveal that mainly olivine fractionation controls the incompatible element budget in Schapenburg komatiites. Moreover, the data indicate that HFSE and REE were not affected by metasomatic processes.

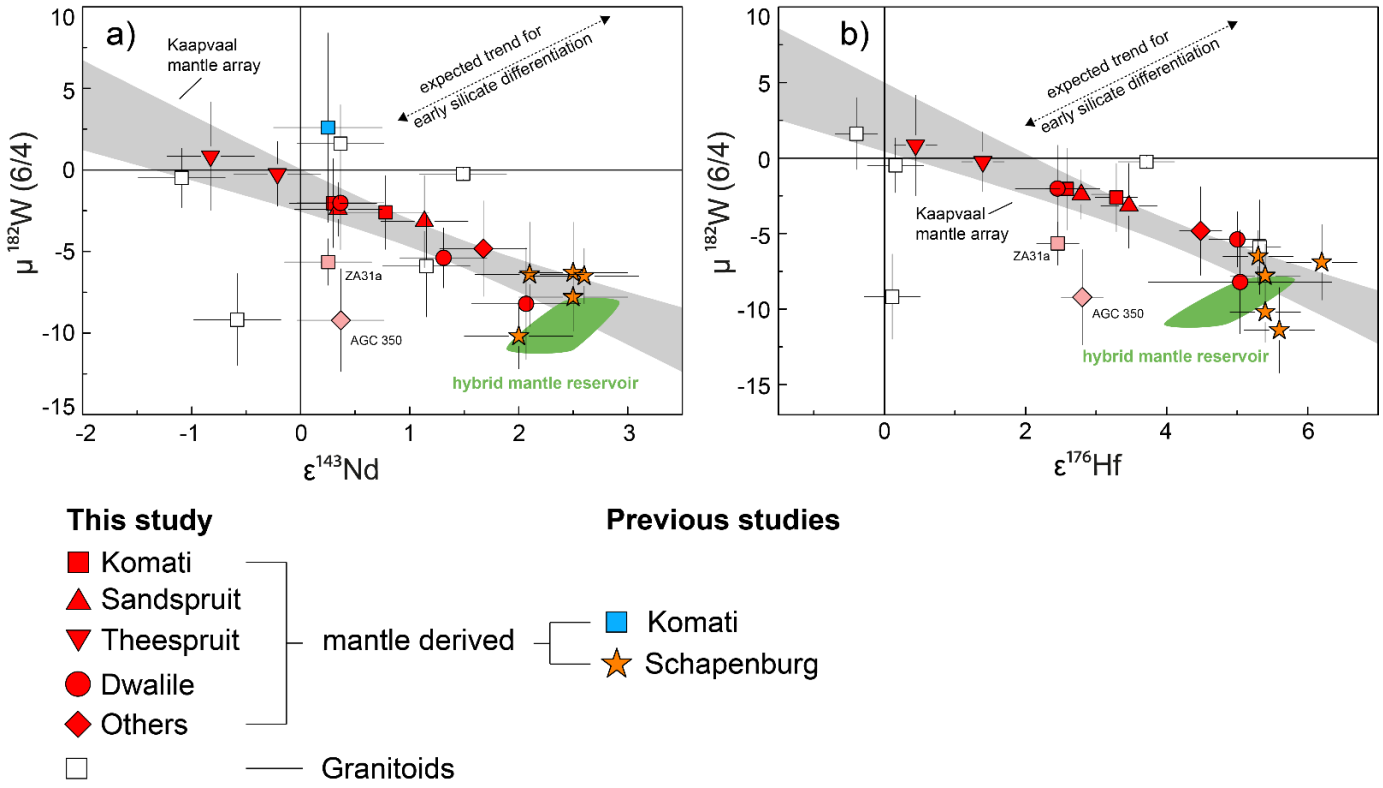

Fig. S4: Measured $\mu^{182}W$ vs. $\varepsilon^{143}Nd_{(t)}$ (a) and $\mu^{182}W$ vs. $\varepsilon^{176}Hf_{(t)}$ (b) for mantle derived and TTG-like mafic rock samples from the Kaapvaal Craton including literature data. The literature data include previously published data for komatiites from the Schapenburg Greenstone Remnant (orange asterisks)(12) and the Komati Formation (blue square)(11, 16). We note that previously published literature data for the Komati Formation only report combined $\mu^{182}W$ vs. $\varepsilon^{143}Nd_{(t)}$ data for one single sample (sample BV 02, blue square)(11,16). The green fields illustrate modeled values of our proposed hybrid reservoir (10-20% restites admixed to ambient depleted mantle). The shaded grey field, referred to as Kaapvaal mantle array, is an uncertainty envelope employing the 95% confidence interval in which of all mantle derived samples are expected to fall. Note, that the negative co-variation displayed by the Kaapvaal mantle array does not follow the expected trend for early silicate differentation (indicated by dashed line in panel b).

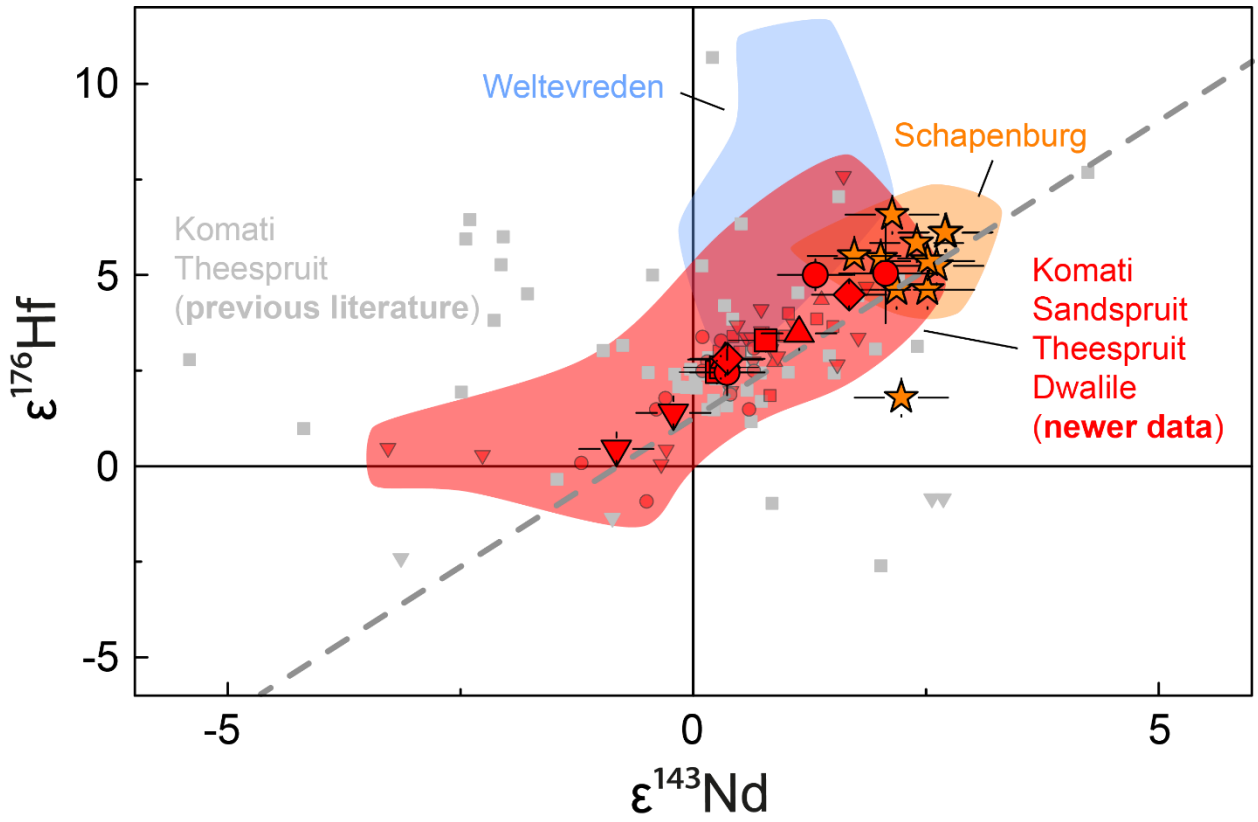

Fig. S5: $\varepsilon^{176}Hf_{(t)}$ vs. $\varepsilon^{143}Nd_{(t)}$ diagram for mantle-derived mafic-ultramafic rocks from this and previous studies. Symbols for our samples are the same as in Fig. 1. Previous literature data from the Komati and Theespruit Formations are displayed by grey symbols(16, 32, 105-107). More recent studies(18, 19) (red symbols) are the data source for samples analyzed in this study (large red symbols). The orange field is defined by komatiites from the Schapenburg Greenstone Remnant (orange stars)(12, 108, 109). The blue field shows the array for the Weltevreden komatiite suite(16, 32). The dashed line shows the Modern Terrestrial Array for MORBs and OIBs ($\varepsilon Hf = 1.55 \times \varepsilon Nd + 1.21$)(33).

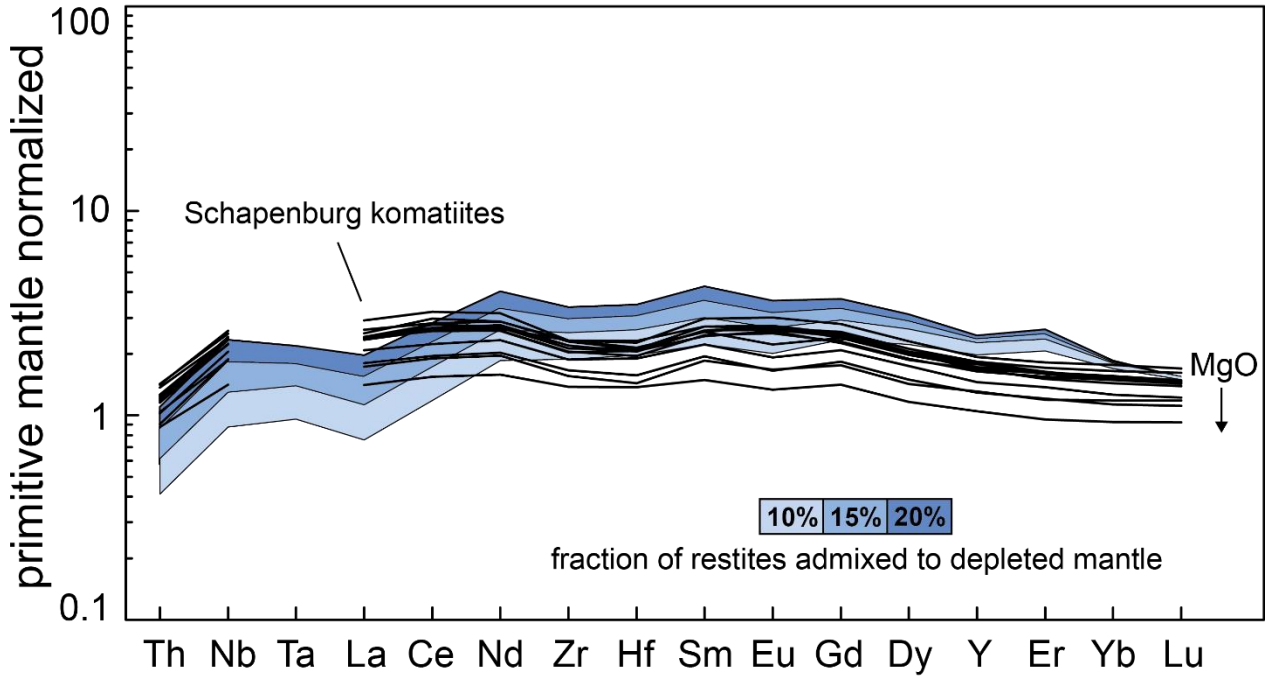

Fig. S6: Incompatibe trace element compositions for komatiites from the Schapenburg Greenstone Remnant (black lines) in comparison to melts generated from our modeled hybrid sources (blue shaded arrays). Data for the SGR komatiites are taken from the literature(12). In our model calculations 10 – 20 % of garnet – rich lower crustal restite admixed to an ambient depleted mantle at 3.55 Ga and subsequent 20 – 30% batch melting of this hybrid source can reproduce the trace element compositions of the SGR komatiites. As outlined previously (12,25) we attribute the variation within the SGR komatiite suite and their more depleted trace element compositions, compared to the modeled patterns, to olivine accumulation as indicated by co-variations between MgO content and incompatible trace element concentrations (Fig. S3).

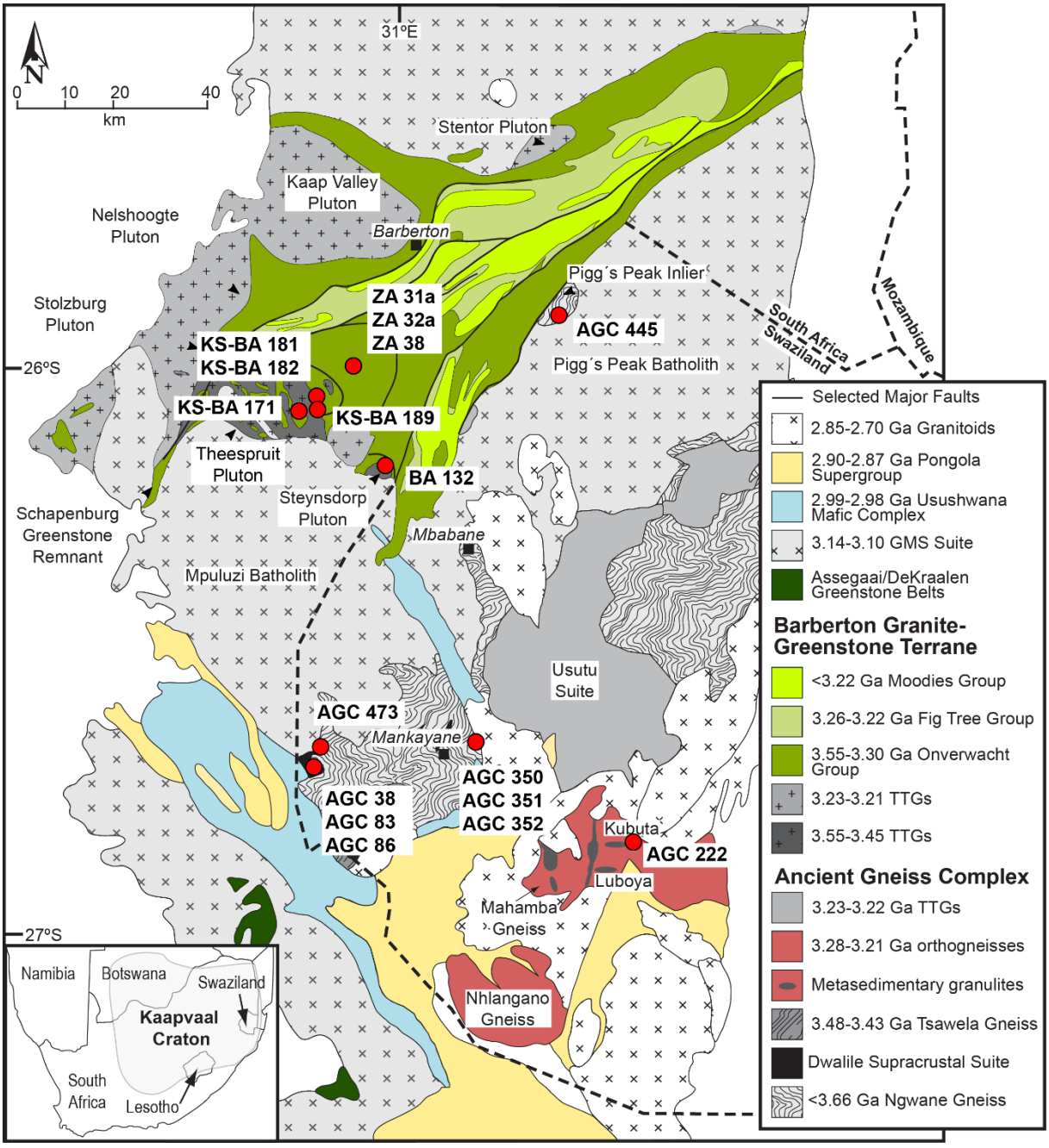

Fig. S7: Simplified geological map of the Kaapvaal Craton, Southern Africa, showing the sample localities covered in this study. The map is taken from Ref. 15 and modified after Ref. 93.

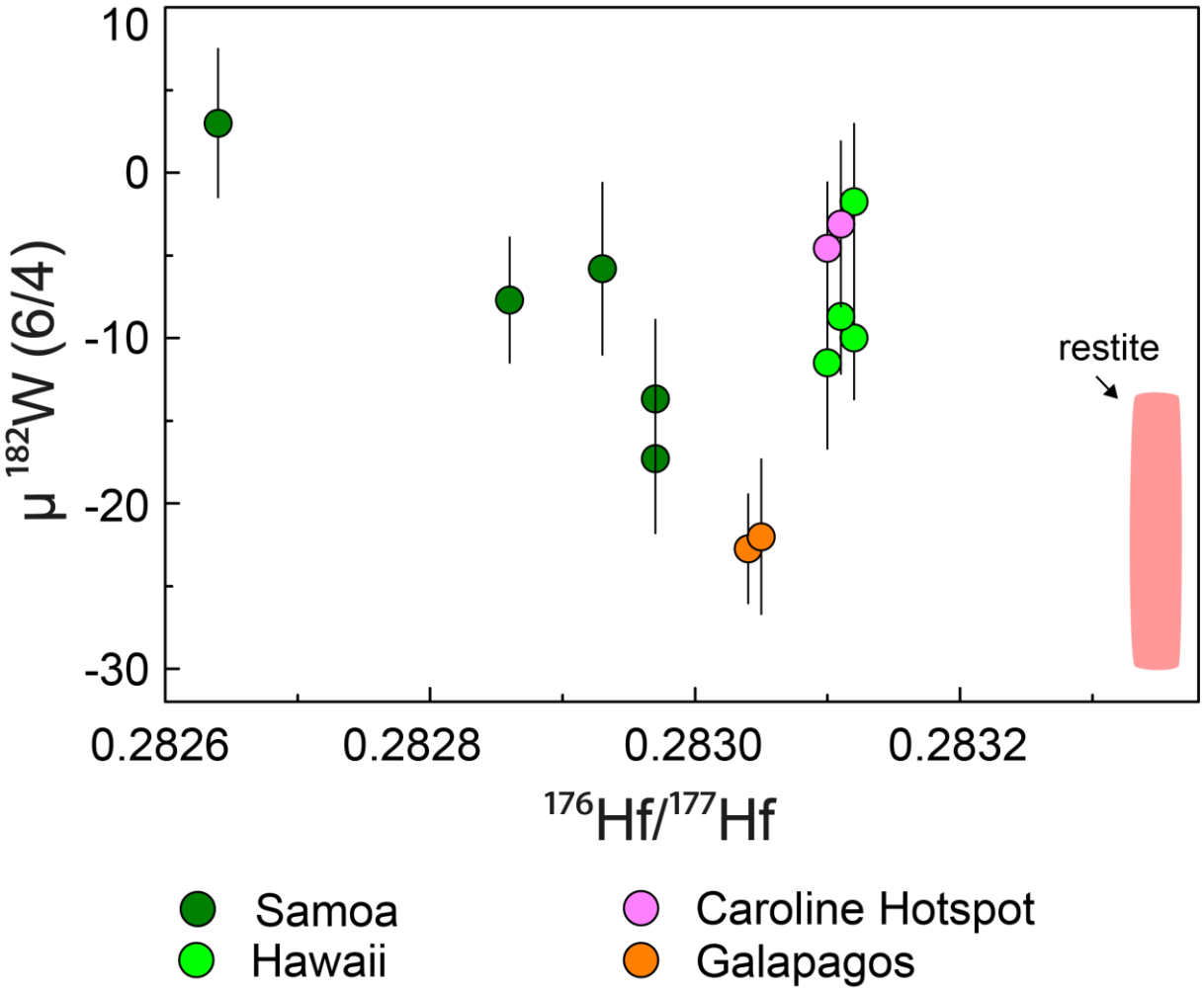

Fig. S8: Compilation of ^{182}W and ^{176}Hf isotope compositions for modern OIBs. Data were compiled from recent studies(8, 48) and references therein. As for $\mu^{182}W$ vs. $^{143}Nd/^{144}Nd$ (see Fig. 5) the global compilation for modern OIBs displays a similar pattern for $\mu^{182}W$ vs. $^{176}Hf/^{177}Hf$, although only less constrained by the limited ^{176}Hf isotope data available. Also shown is the present ^{182}W and ^{143}Nd isotope composition calculated for restites that remained after prolonged TTG formation (4.35 – 4.25 Ga) via partial anataxis of a mafic protocrust that formed between 40 and 50 Ma after solar system formation (pink array).

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